Interludes

In the **two interludes**, you will explore **two topics** that are outside of the traditional purview of astronomy, physics, and the physical sciences, as well as mathematics and computer science.

These interludes reach out to the "Arts, Letters. and Values" (ALV) and "Cultures, Societies, and Individuals" (CSI) knowledge domains.

You will write **two short papers**, one for each interlude, and make **one team presentation** for *interlude I* or for *interlude II*.

2x Papers

- You will write <u>two</u> short papers, one for each interlude.
- The papers should be <u>5 pages long</u> (double spaced).
- **Each paper must include at least one figure** (i.e. photo, diagram, image, plot, table).
- In the case of ***interlude I***, this <u>figure must be of your own making</u>, and it should be used to explain and support the arguments and information in your paper (i.e. it is not decoration); you can also have additional figures that are taken from other sources, so long as they are properly referenced in your bibliography (note: these additional figures are not included in the minimum 5 page count).
- In the case of ***interlude II***, this figure can of your own making or from another source (and properly referenced in the bibliography). Each student must turn their own distinct paper.
- **Format:** 12 point, Times New Roman, 1" margins, 8" × 11" paper hardcopy.

Team Presentation

- You will participate in a **team of four students** to study, explore, and develop an interlude topic.
- Your team will give a short **10 minute presentation** on one of the topics.
- You only need to make one presentation (either for interlude I or for interlude II).

Note: Your paper is expected to be on the same topic as your team presentation (though, if you really want to explore another topic you can). The paper is <u>not</u> a team effort.

Interlude 1
Humanity in the Solar System

Presentations: November 4-6.
Paper deadline: November 8.

Interlude 2
Space Art

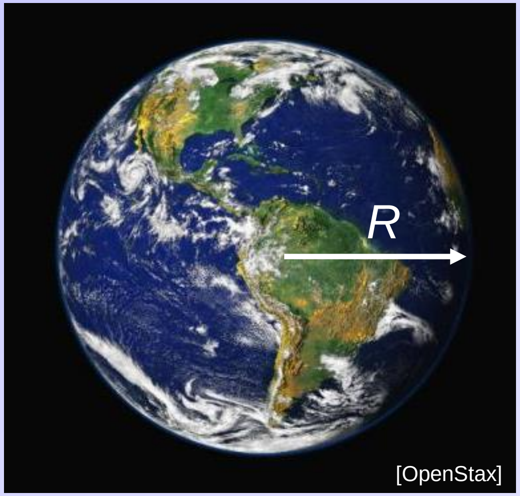
Presentations: December 2-4.
Paper deadline: December 6.

Today's Topics

Friday, October 4, 2019 (Week 5, lecture 15) – Chapter 7.

1. Density of Planets
2. Formation of the Solar System
3. Age of the Solar System

Planet Density



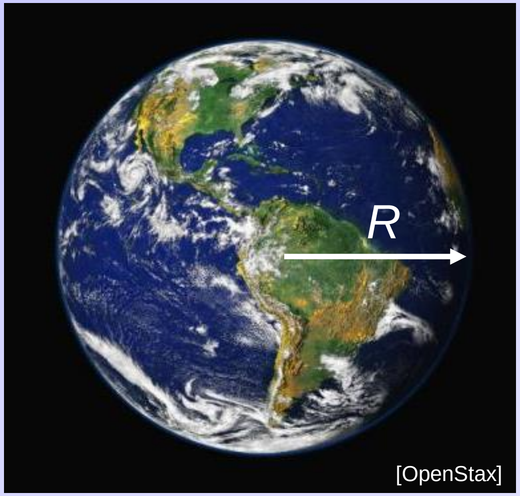
Q: How do you calculate density ?

Answer:

Density = $\rho = \frac{Mass}{Volume} = \frac{Mass\ of\ Planet}{Volume\ of\ Planet}$

$Volume\ of\ a\ Sphere = V_{sphere} = \frac{4}{3}\pi R^3$ with $R = radius\ of\ sphere/planet$

Planet Density



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Densities of planetary materials

water/ice H_2O = 1 g/cm^3

liquid hydrogen = 0.07 g/cm^3

liquid helium = 0.1 g/cm^3

liquid nitrogen = 0.8 g/cm^3

liquid methane = 0.4 g/cm^3

solid CO_2 = 1.6 g/cm^3

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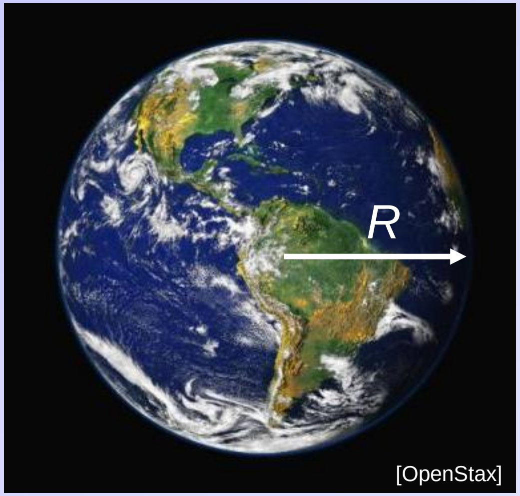

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limestone ~ 2.6 g/cm^3

granite ~ 2.7 g/cm^3

basalt ~ 3.0 g/cm^3

rock

iron ~ 9 g/cm^3

nickel ~ 9 g/cm^3

uranium ~ 19 g/cm^3

iridium ~ 22.7 g/cm^3

Composition of Planets

water/ice H_2O = 1 g/cm^3
liquid hydrogen = 0.07 g/cm^3
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rock

iron ~ 9 g/cm^3
nickel ~ 9 g/cm^3
uranium ~ 19 g/cm^3
iridium ~ 22.7 g/cm^3

Name	Distance from Sun (AU)[2]	Revolution Period (y)	Diameter (km)	Mass (10^{23} kg)	Density (g/cm^3)[3]
Mercury	0.39	0.24	4,878	3.3	5.4
Venus	0.72	0.62	12,120	48.7	5.2
Earth	1.00	1.00	12,756	59.8	5.5
Mars	1.52	1.88	6,787	6.4	3.9
Jupiter	5.20	11.86	142,984	18,991	1.3
Saturn	9.54	29.46	120,536	5686	0.7
Uranus	19.18	84.07	51,118	866	1.3
Neptune	30.06	164.82	49,660	1030	1.6

rocks + metals (Mercury, Venus, Earth, Mars)

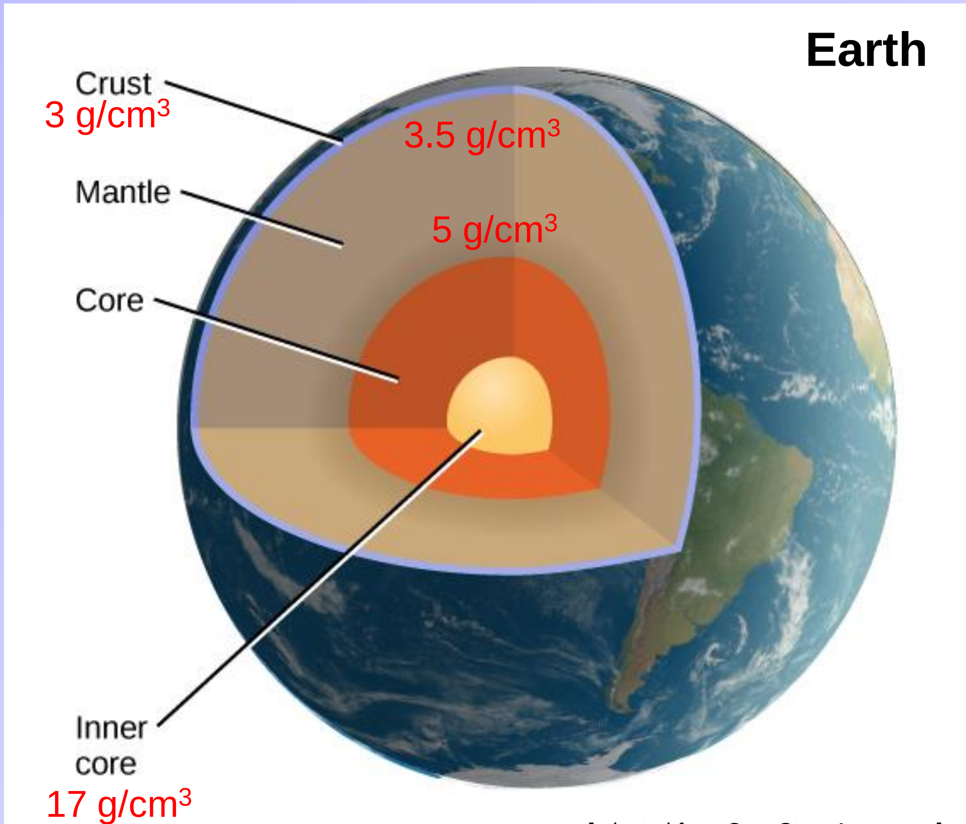
icy (Jupiter, Saturn, Uranus, Neptune)

[OpenStax: Astronomy]

Differentiation: Planet Density is Not Uniform

- The density of a planet is largest at the center and weakest at the surface.
- **Differentiation:** the composition of a planet varies with depth.

  → Denser materials/elements sank under gravity, when planet was a mix of hot liquids (or gases).



[adapted from OpenStax: Astronomy]

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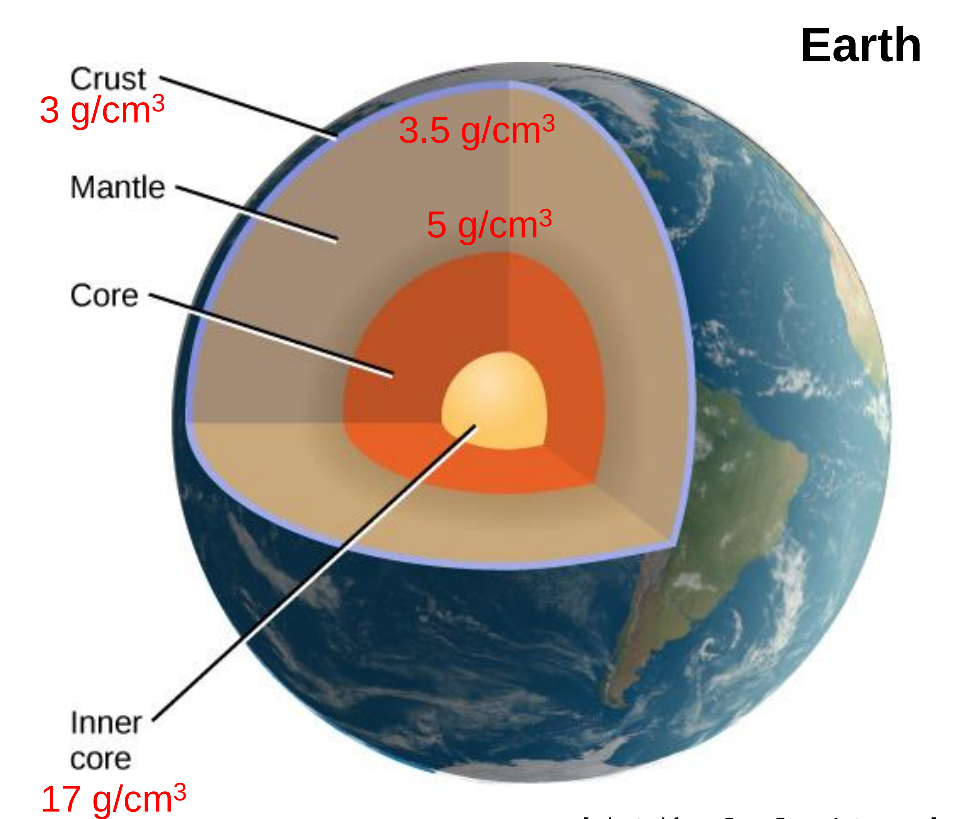
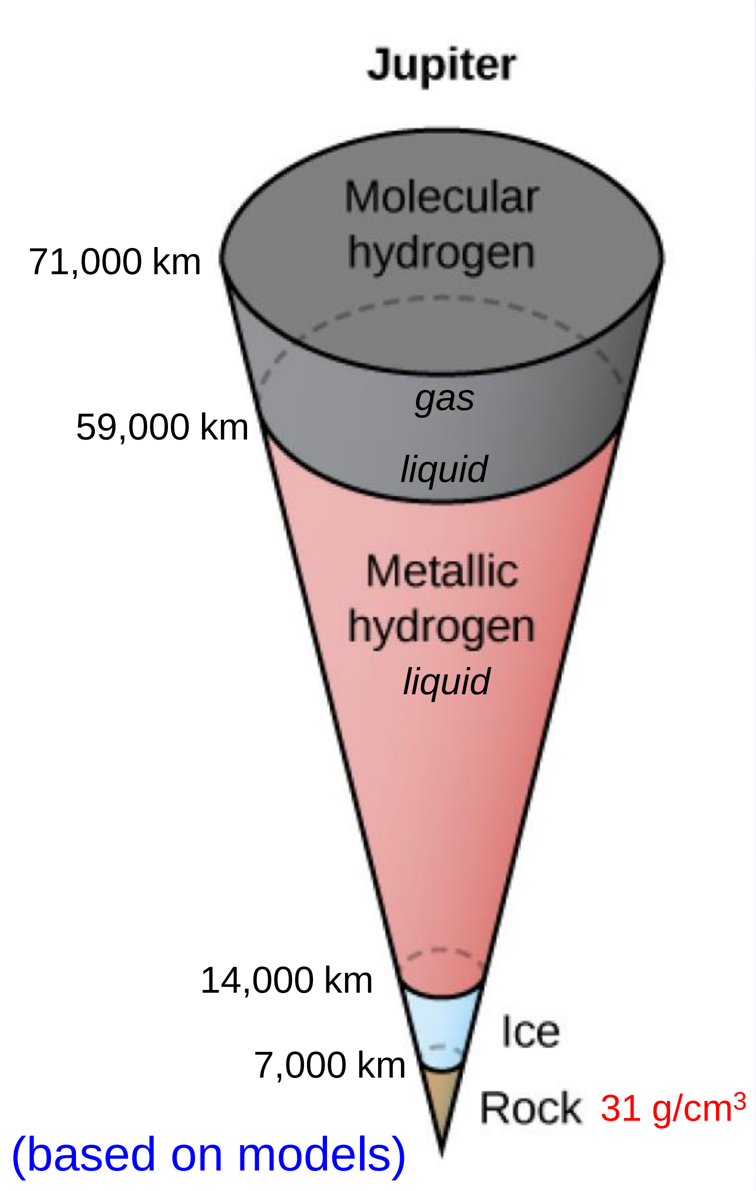


[adapted from OpenStax: Astronomy]



[adapted from OpenStax: Astronomy]

Differentiation: structure of solar system

Solar Nebula

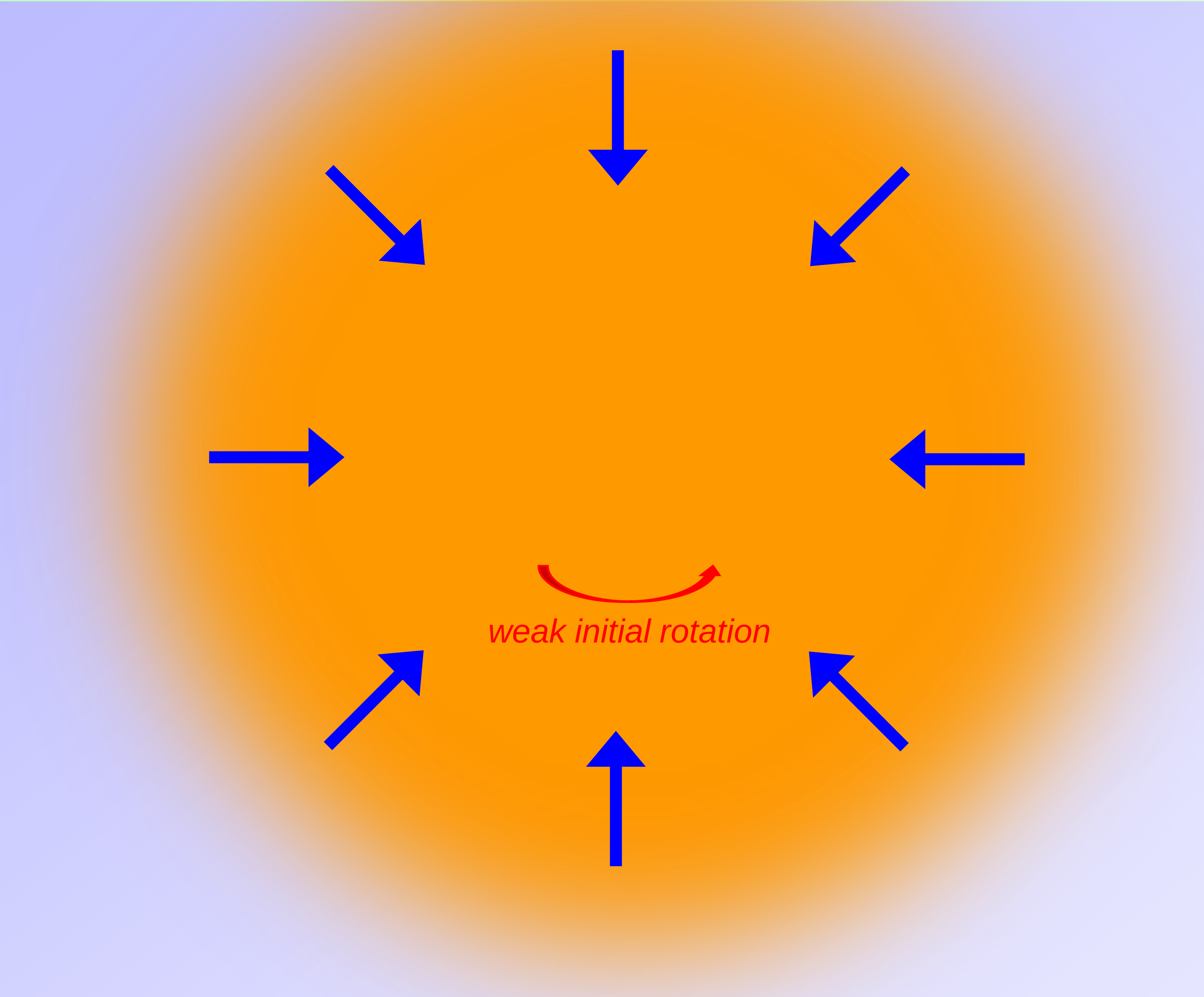
This artist's conception shows the flattened cloud of gas and dust from which the Solar System formed [OpenStax: Astronomy].

- Icy and rocky planetesimals (precursors of the planets) can be seen in the foreground.
- The bright center is where the Sun is forming.


[William K. Hartmann, Planetary Science Institute]

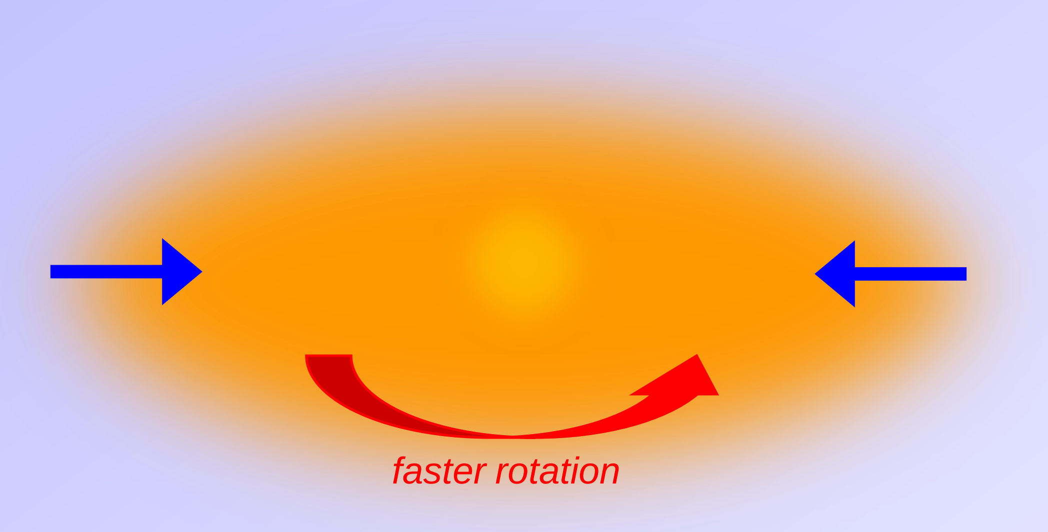
- **Mutual gravity** pulls dust, particles, material, and gas inwards.
- **Contraction:** As the solar nebula contracts it **heats up** (energy conservation), spins faster (angular momentum conservation), and flattens out.
- **Condensation:** As the nebula cools (blackbody radiation) heavy element gases condense around dust particles. Hydrogen and helium do not condense.
- **Accretion of planetesimals:** Solid particles collide and stick together to progressively start planets. The central region gets dense enough to **ignite fusion**.
- **Differentiation:** Hydrogen based molecules can condense far from the center (where it is colder), but not near the center where it is hotter. Heavier elements can condense closer to the Sun where it is hotter.

Step 0: large cloud of gas & dust



Mutual gravity pulls gas, dust, particles, and material inwards.

Step 1: Contraction



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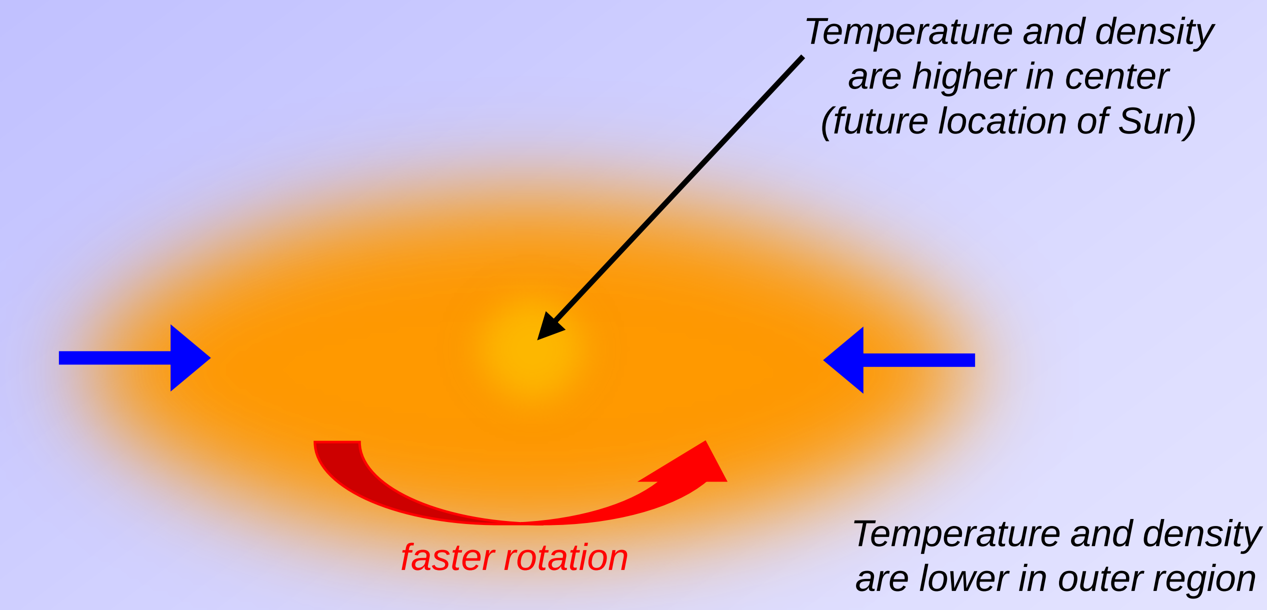
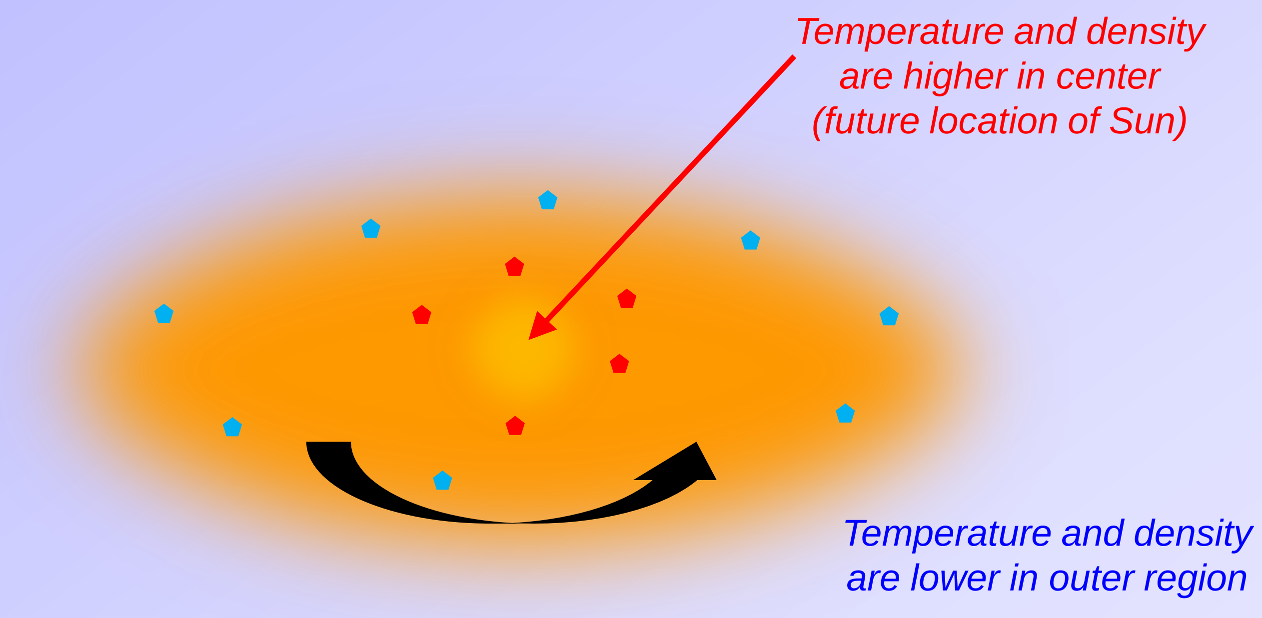


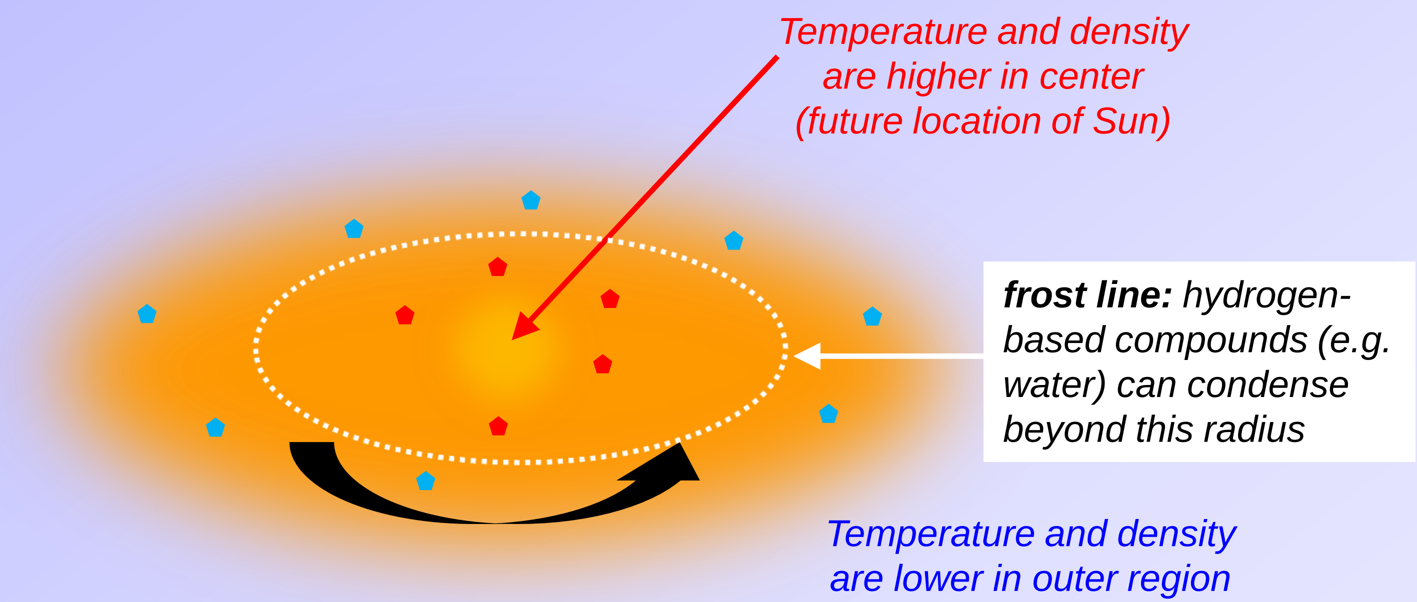
As the solar nebula contracts it **heats up** (energy conservation), spins faster (angular momentum conservation), and flattens out.

Step 2: Condensation



As the nebula cools (blackbody radiation) <u>heavy element gases</u> condense around dust particles. Hydrogen and helium do not condense, but hydrogen-based molecules can in the cooler outer parts.

Step 2: Condensation – "frost line"



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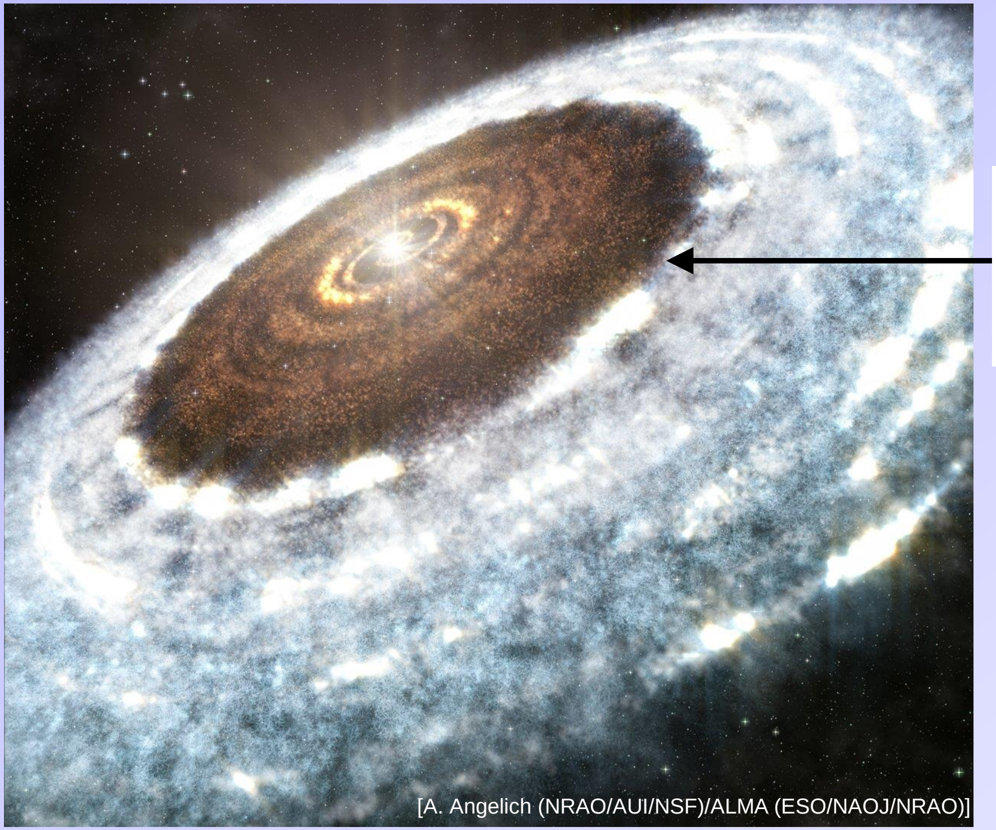
Artist's impression of the water snowline around the star V883 Orionis [Wikipedia].



***frost line:** hydrogen-based compounds (e.g. water) can condense beyond this radius*

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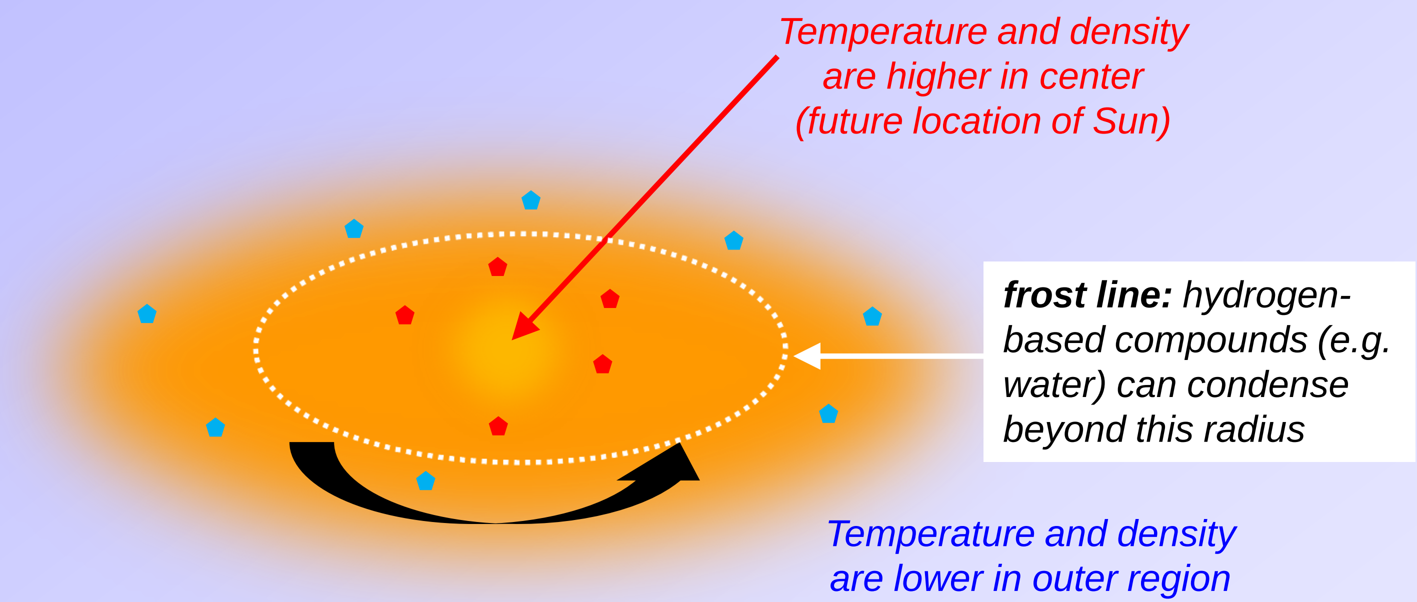

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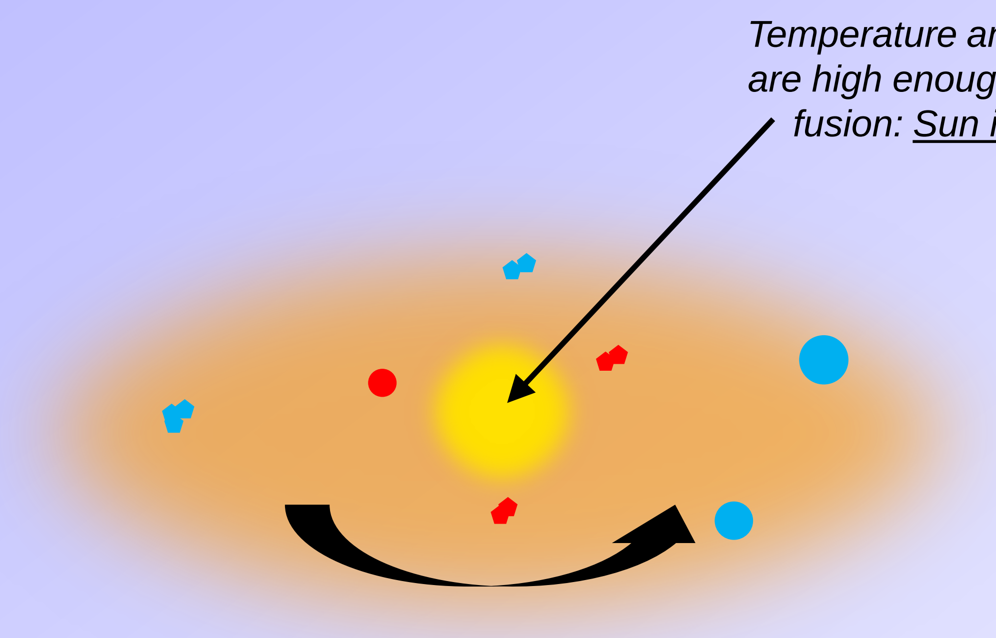
Step 3: Accretion of Planetesimals

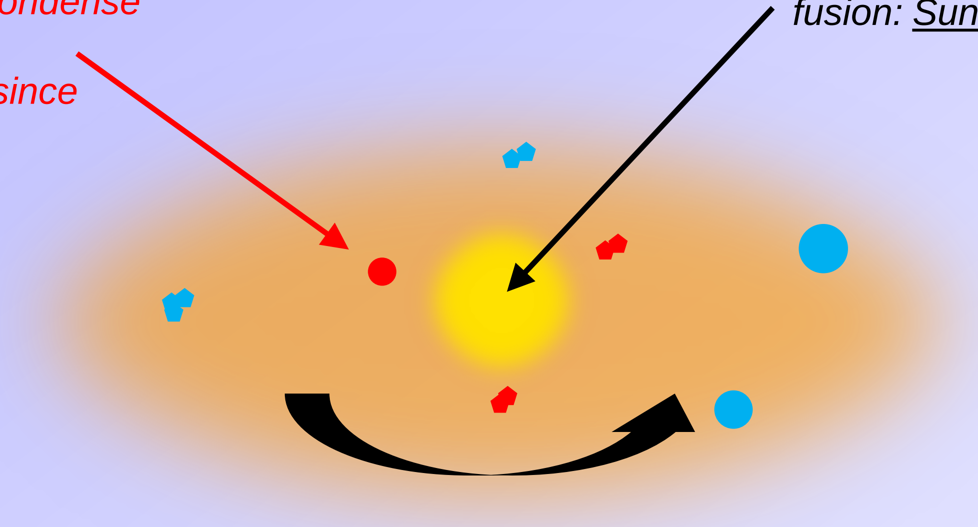


- **Accretion of planetesimals:** Solid particles collide and stick together to progressively start planets. Their gravity becomes strong enough to collect gases.
- **Star ignition:** The central region gets dense enough to **ignite fusion**.

Step 3: Accretion of Planetesimals

Inner planetesimal tend to be richer in heavy elements, which condense at higher temperature. They tend to be smaller, since they sweep a smaller area (gathering less material).

Temperature and density are high enough to ignite fusion: <u>Sun is born</u>.



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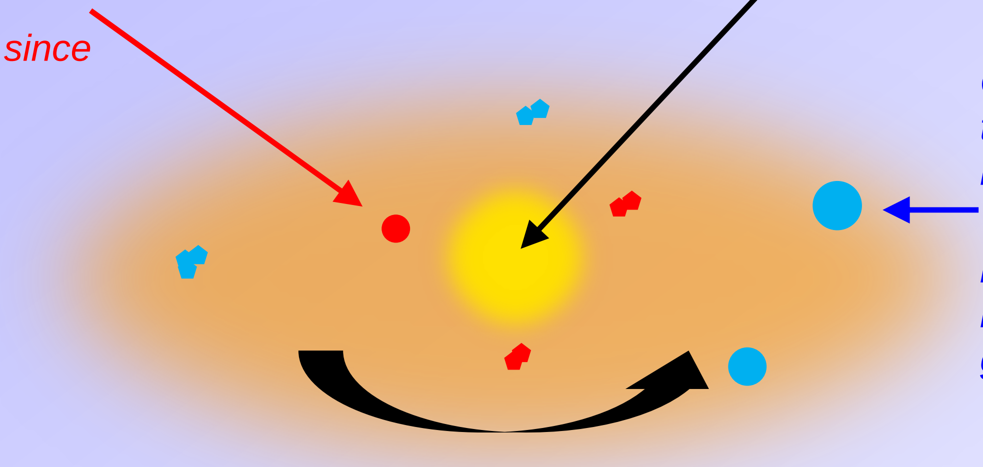
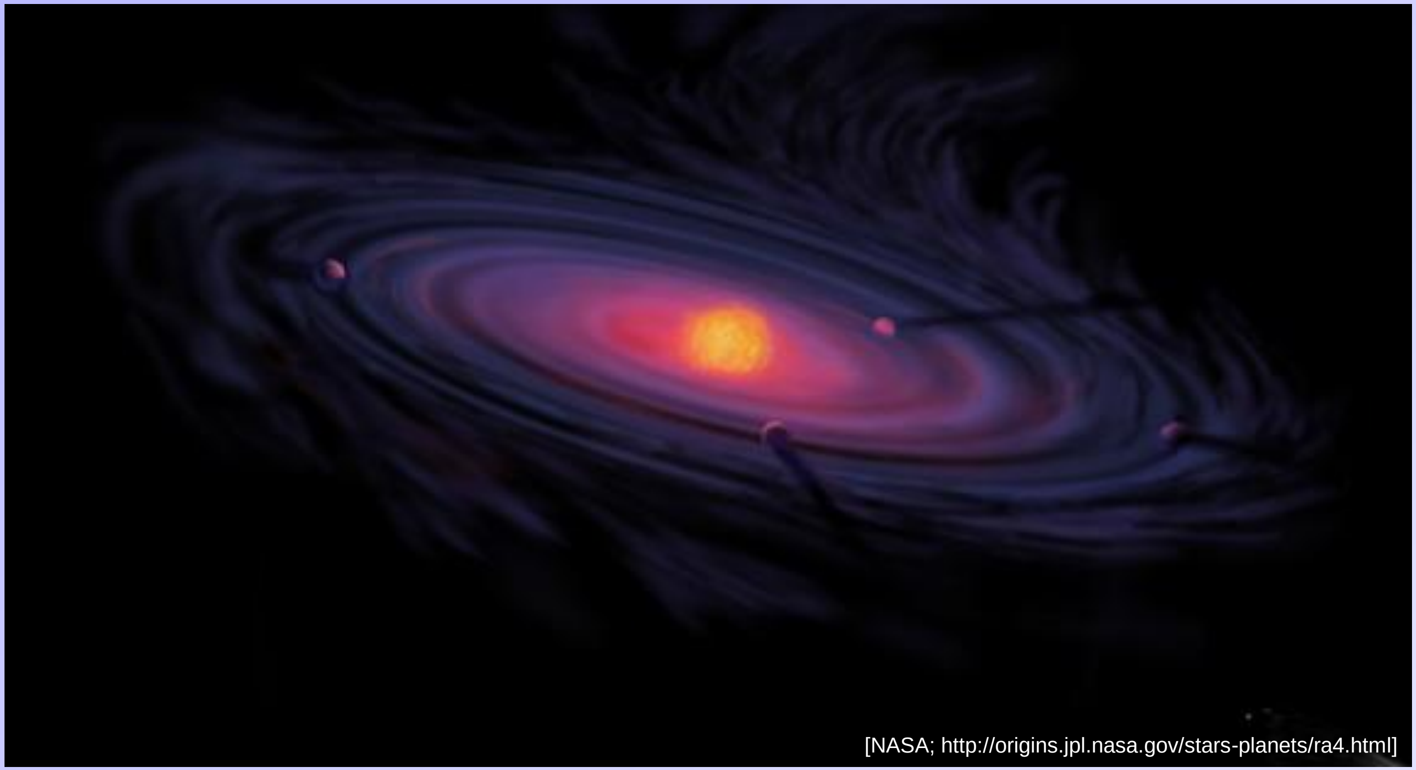


Outer planetesimals tend to be more icy and hydrogen rich.
They tend to be bigger because they sweep out a larger region, so they can gather more material.

- **Accretion of planetesimals:** Solid particles collide and stick together to progressively start planets. Their gravity becomes strong enough to collect gases.
- **Star ignition:** The central region gets dense enough to **ignite fusion**.

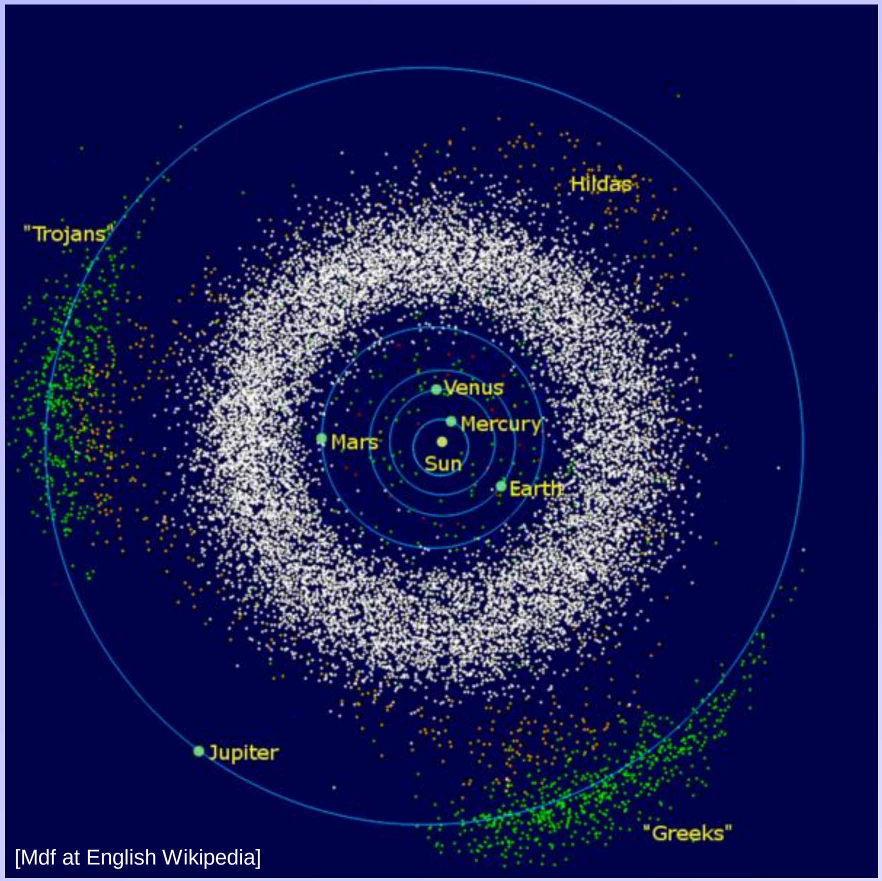
Step 3: Planetesimals to Planets



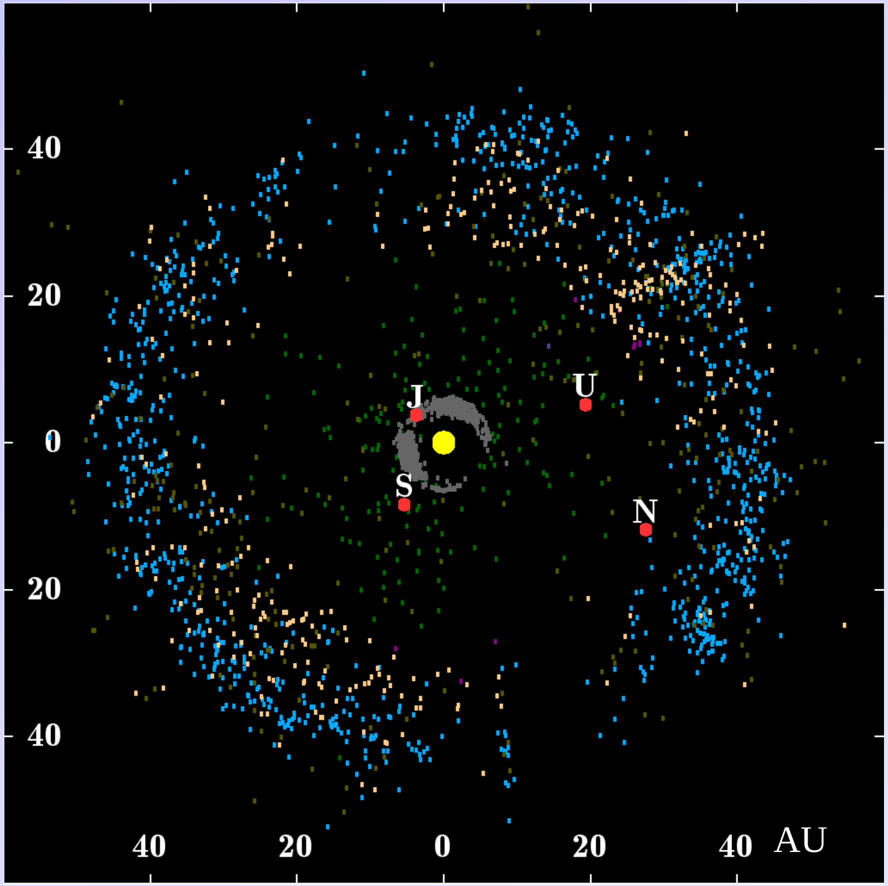
- As the planetesimal collide and stick together, they become bigger and evolve into planets. In doing so, they clear out their orbits.
- Near circular orbits are more stable, since more eccentric elliptical ones can lead to collisions between planetesimals/planets.

Solar System Planets + Planetismals

Inner Solar System + Jupiter



Outer Solar System with Gas Giants



[By WilyD at English Wikipedia, CC BY-SA 3.0]

Asteroid (white, green, brown) are left over planetesimals.

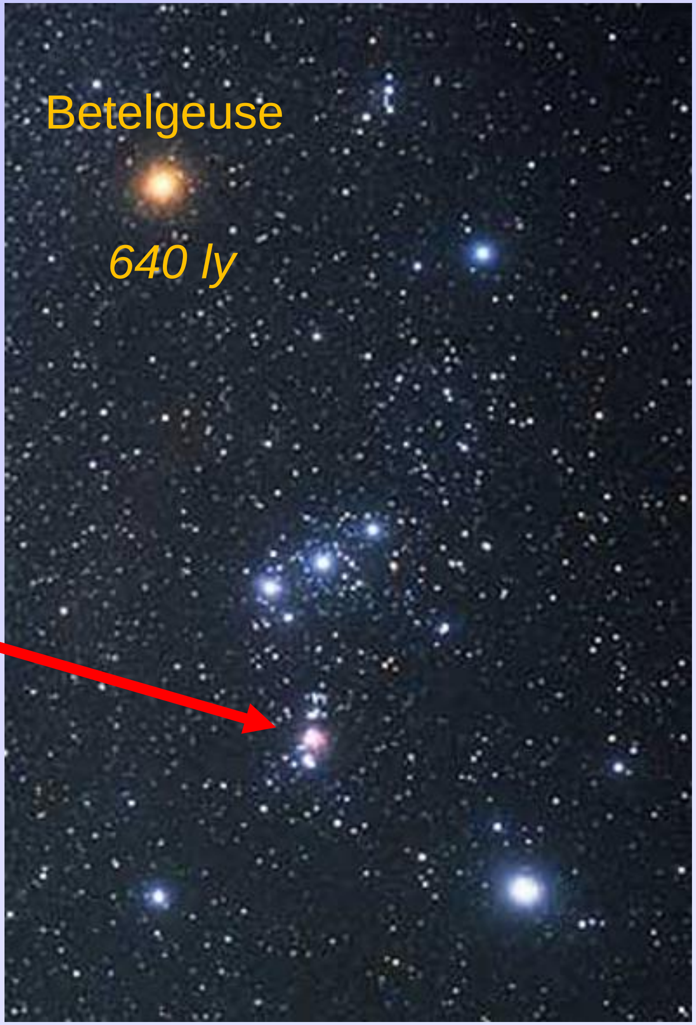
Kuiper belt objects (blue, beige, green) are icy left over planetesimals in the region of the gas giants and beyond.

Nascent Protoplanetary Systems

Constellation: **Orion**

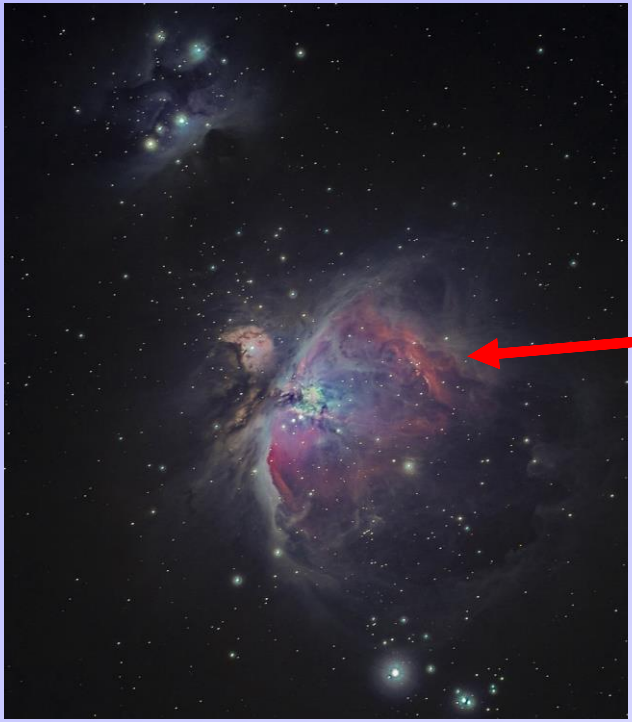
Nascent Protoplanetary Systems

Orion Nebula

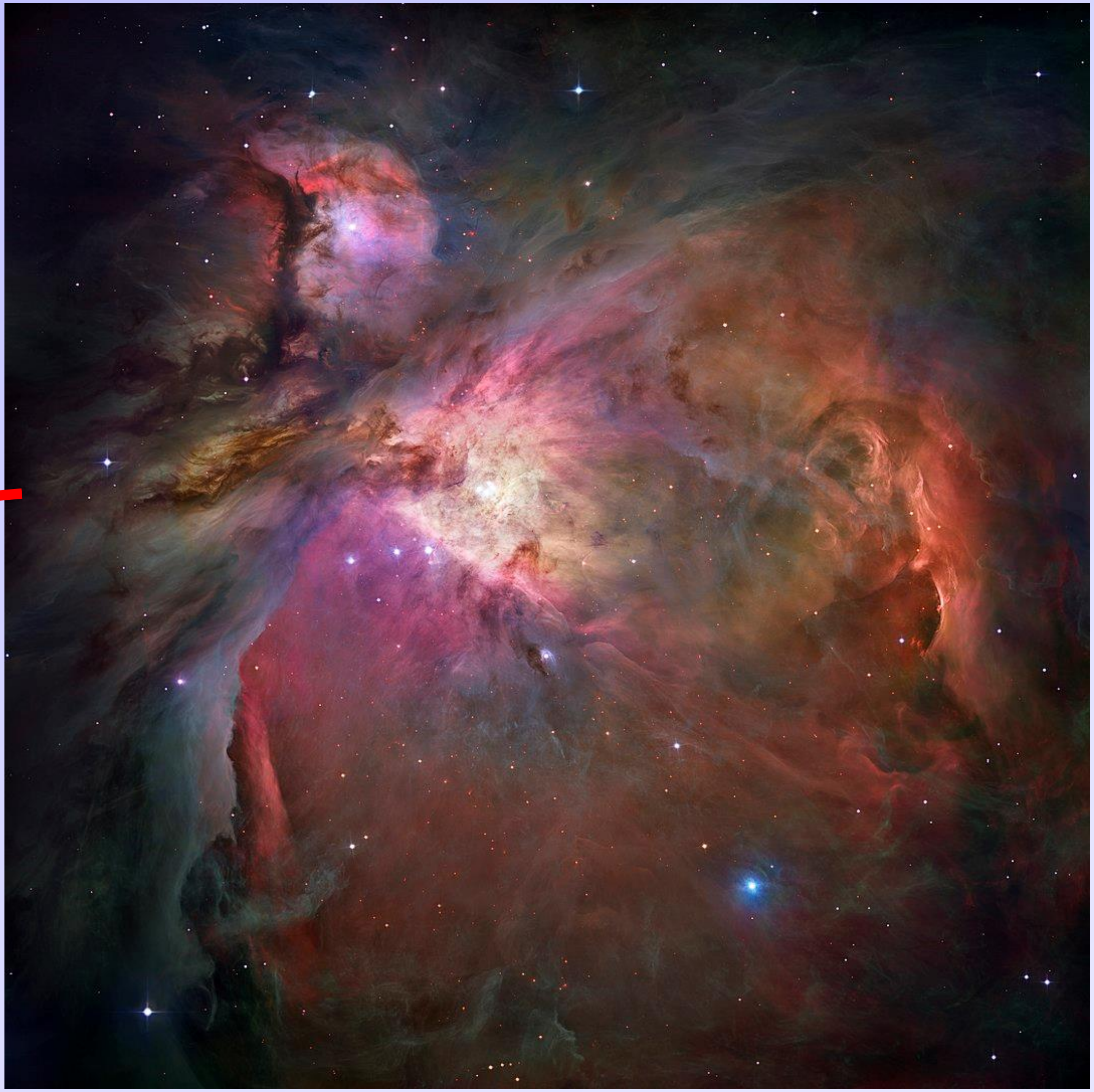


Constellation: **Orion**

Nascent Protoplanetary Systems

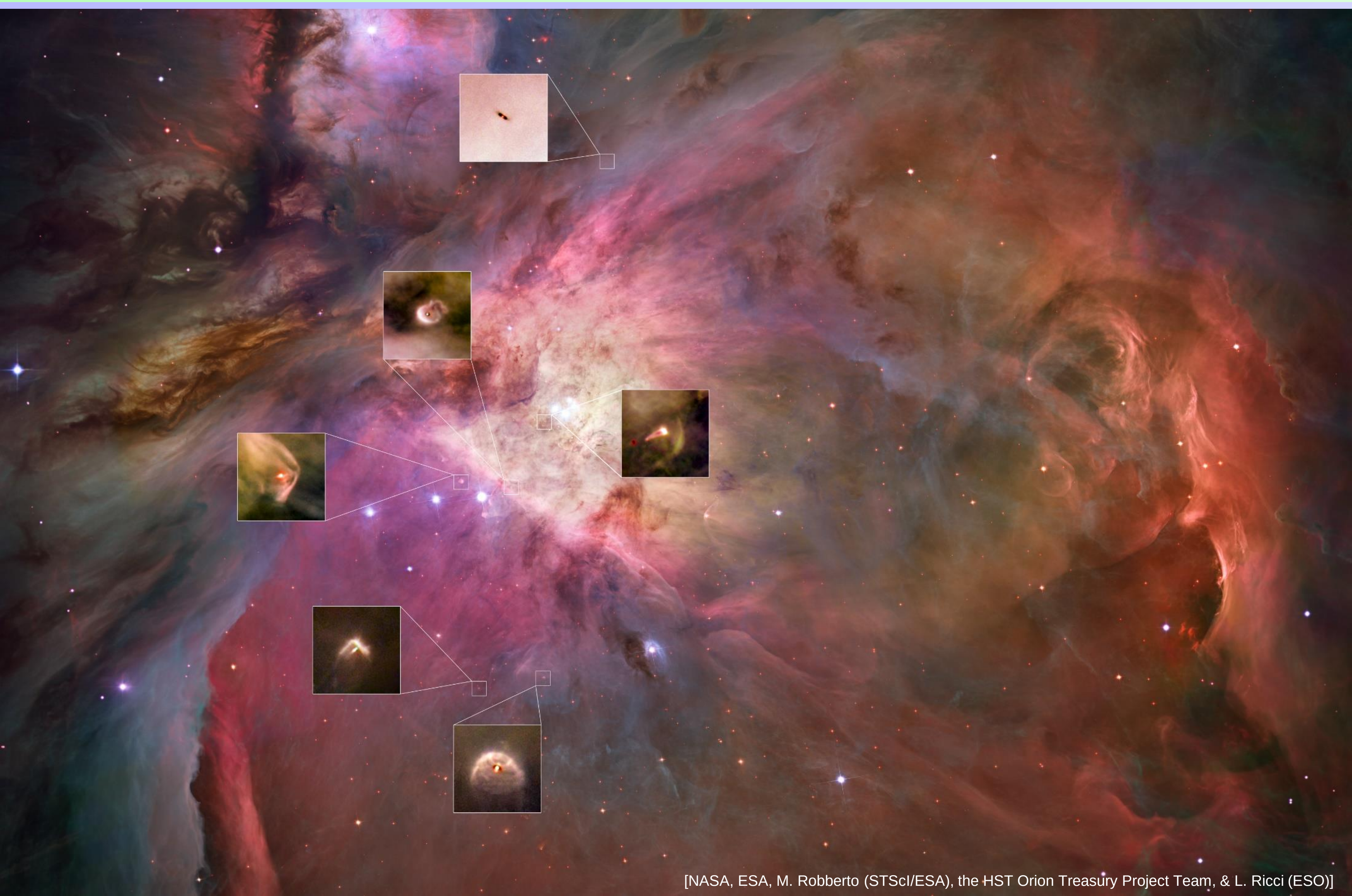


Orion Nebula



[NASA, ESA, M. Robberto (Space Telescope Science Institute/ESA) and the Hubble Space Telescope Orion Treasury Project Team]

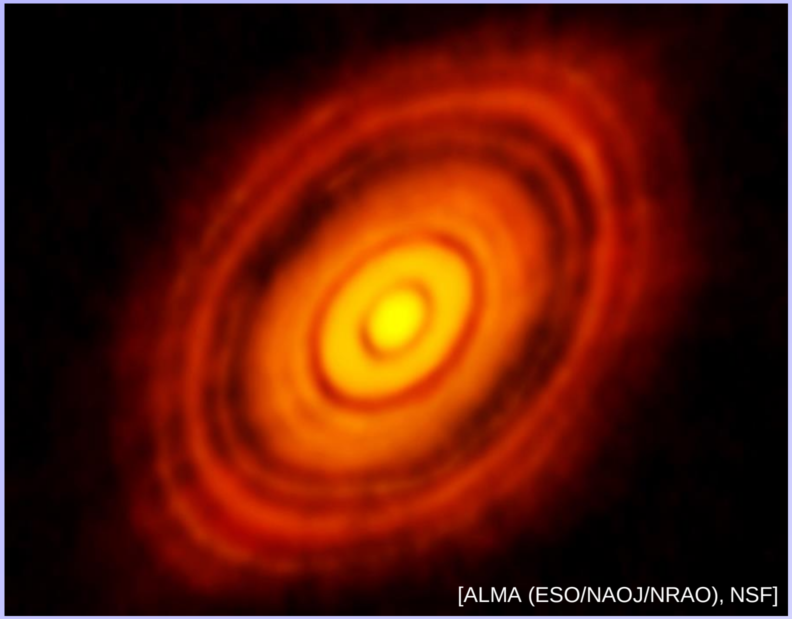
Nascent Protoplanetary Systems



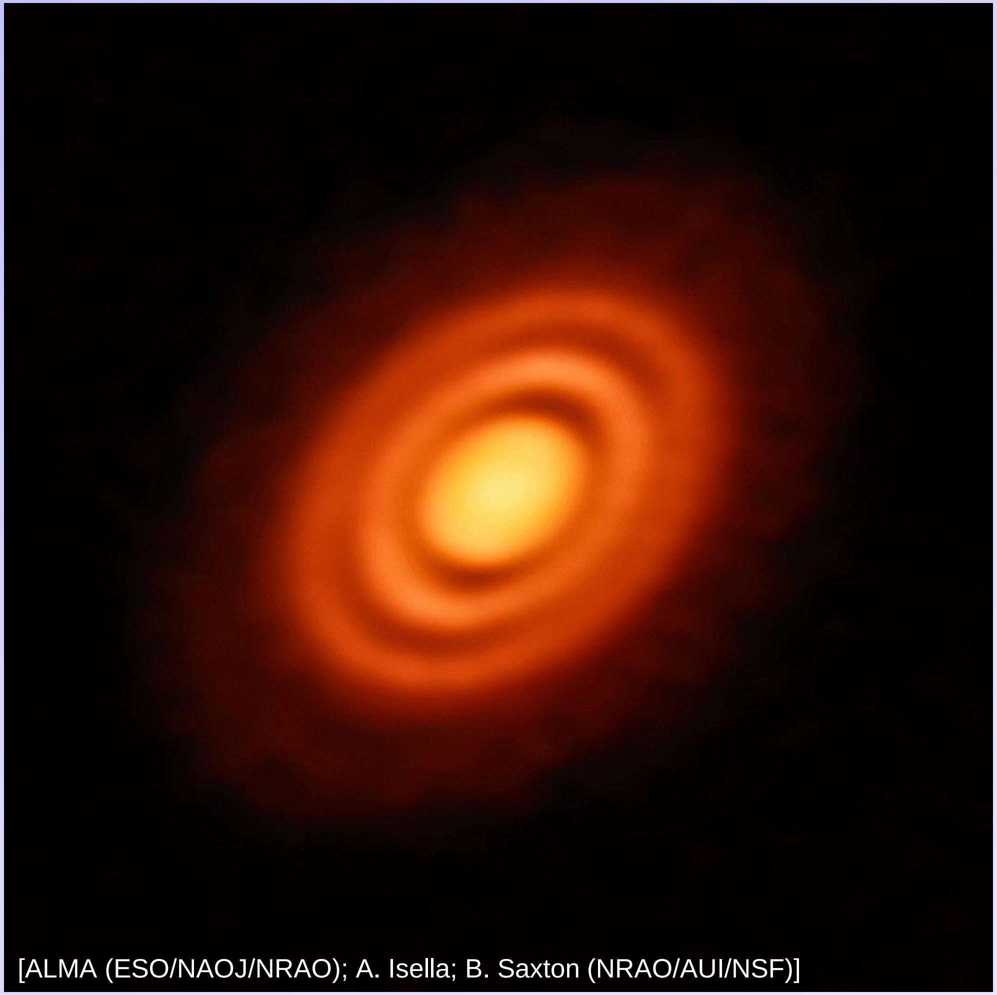
[NASA, ESA, M. Robberto (STScI/ESA), the HST Orion Treasury Project Team, & L. Ricci (ESO)]

Nascent Protoplanetary Systems

Protoplanetary Disks – mm wave

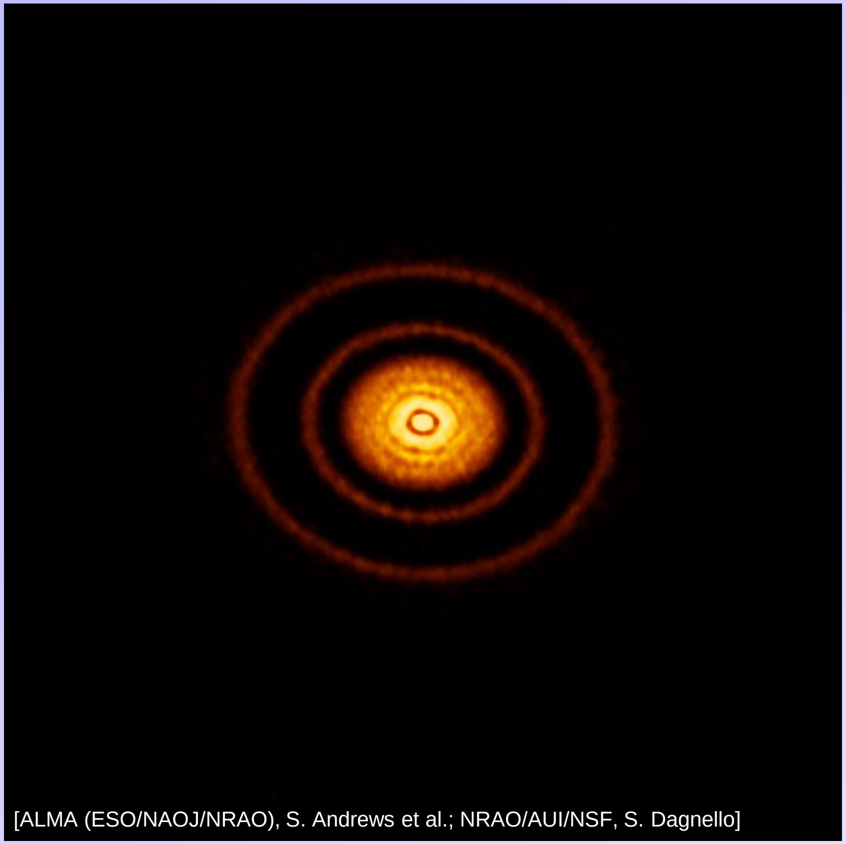


The Protoplanetary Disk of the young star HL Tauri
(in Milky Way galaxy, Taurus constellation)



Cloud of gas and dust surrounding the young star HD 163296.
(in Milky Way galaxy, Sagittarius constellation)

Protoplanetary Disks – mm wave



Protoplanetary disk around the young star AS 209.

(in Milky Way galaxy, Ophiuchus constellation)

How Old is the Solar System ?