

# Today's Topics

Wednesday, March 19, 2025 (Week 7, lecture 18) – Chapters 22, 23.

## Sun-like stars: old age to death

- A. Main sequence to red giant.
- B. Planetary Nebulae.
- C. White Dwarfs.

**Problem Set #6** is available on ExpertTA and is due on Friday, March 21 by 9:00 am

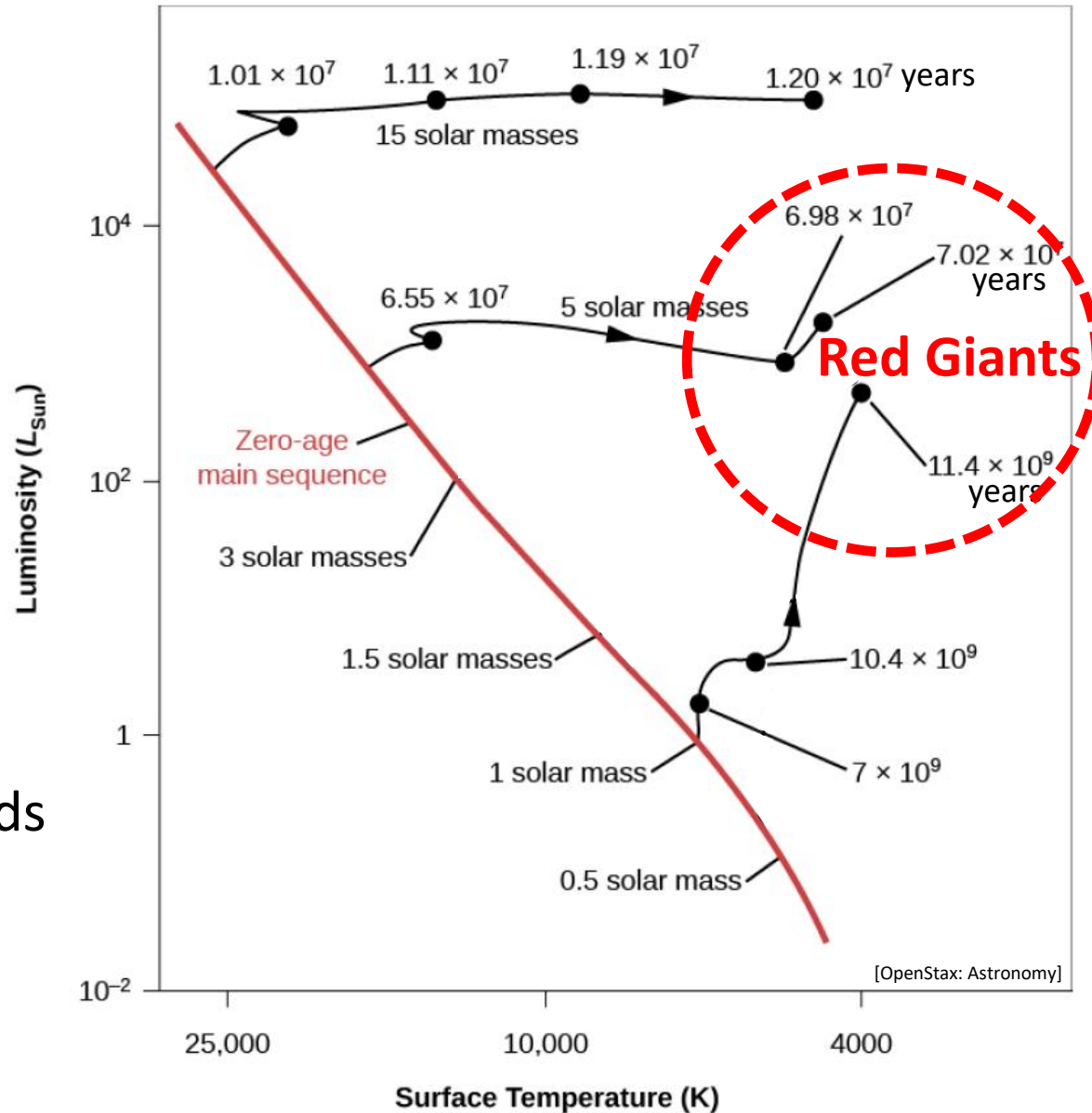
# Stellar Evolution: on the H-R Diagram

## Light stars (sun-like & smaller)

- Yellow and red color.
- cooler and dimmer.
- Long lived.  
→ > 10 billion years.

## Old age

- Stars evolve quickly towards the upper right corner.  
→ More luminous, but cooler.



# Evolution of Sun-like Stars

Stage	Time in This Stage (years)	Surface Temperature (K)	Luminosity ( $L_{\text{Sun}}$ )	Diameter (Sun = 1)
Main sequence	11 billion	6000	1	1
Becomes red giant	1.3 billion	3100 at minimum	2300 at maximum	165
Helium fusion	100 million	4800	50	10
Giant again	20 million	3100	5200	180

white dwarf  
(+ planetary nebula)

“forever”

40,000 K  $\rightarrow$  4,000 K

$\sim 1 \rightarrow 0.01$

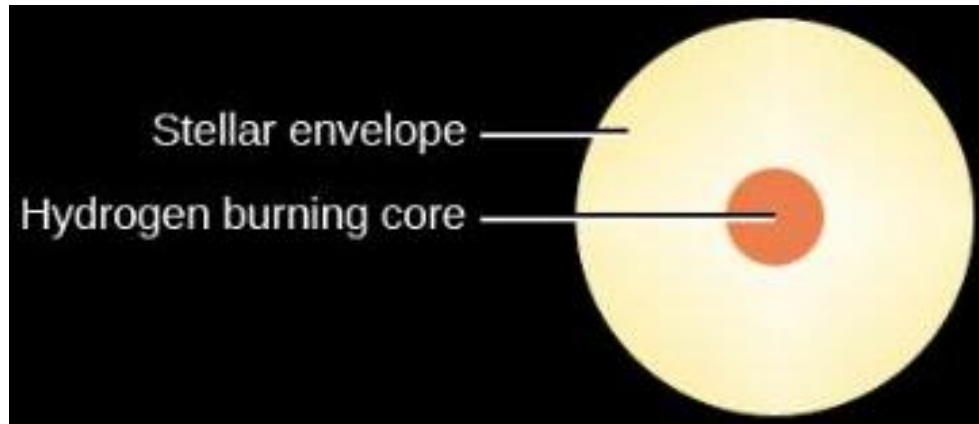
$\sim 0.01$

# Becoming a Red Giant

## 1. Main sequence operation

Proton-proton fusion chain in core

*4x Hydrogen → 1 helium*



# Becoming a Red Giant

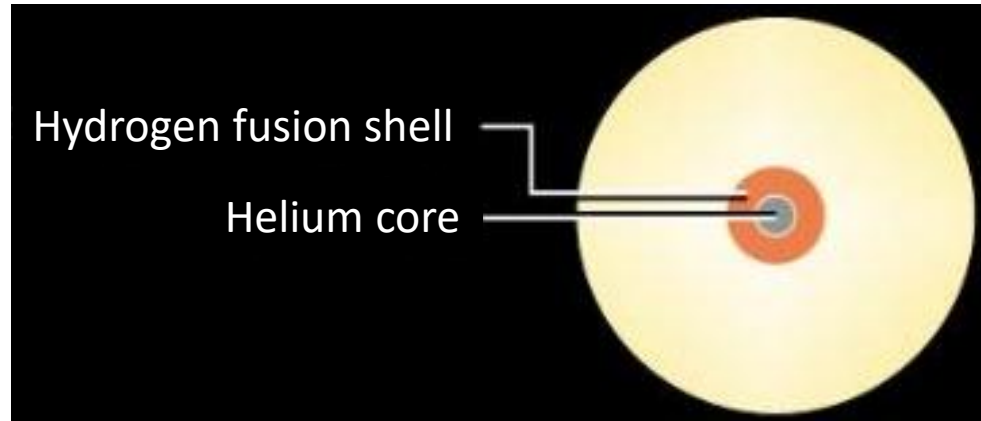
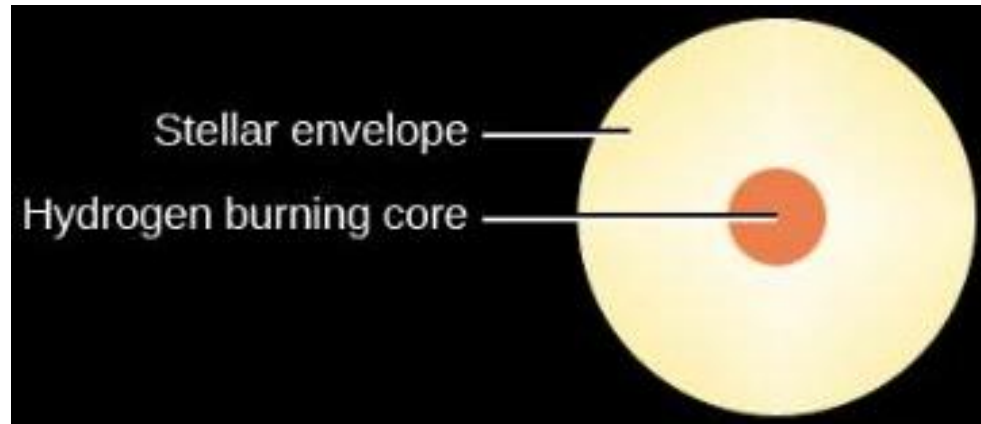
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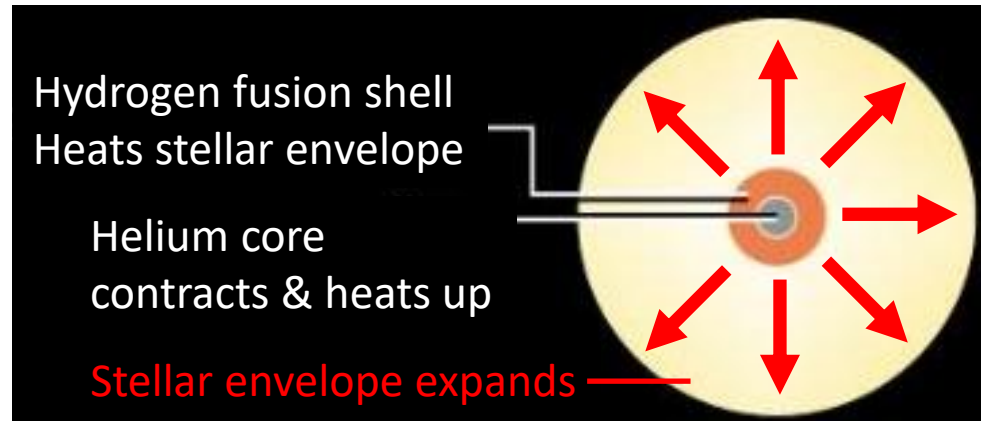
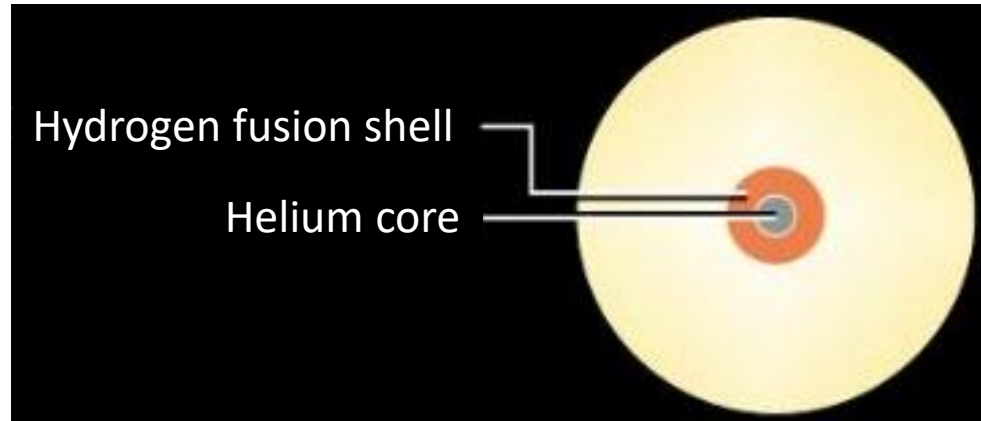
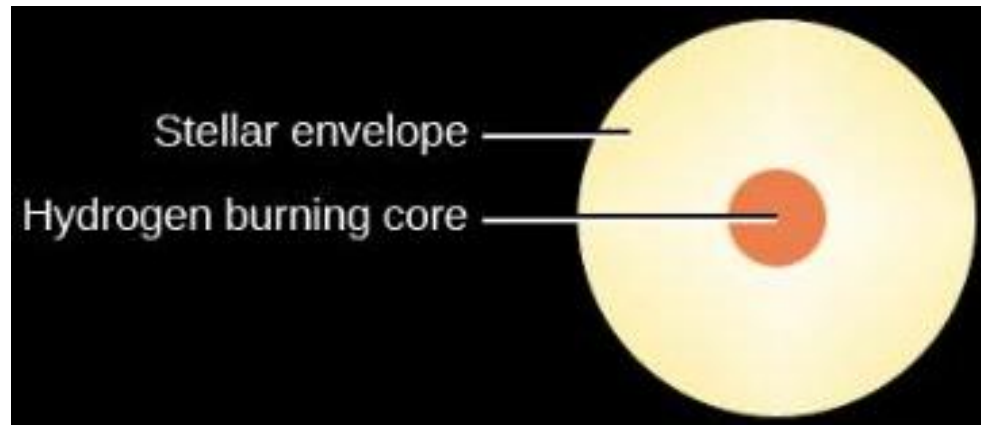


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## 3. Expansion to red giant

- **Heat** from new hydrogen shell fusion is **significant** and heats up outer hydrogen in stellar envelope.
- Stellar envelope heats up and **expands** (outer layer then cools).
- **Helium core continues to contract and heat up**.

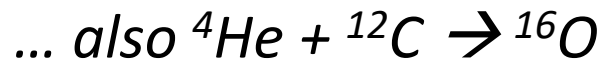


# Helium Fusion

- At  $T \approx 100,000,000$  K, helium nuclei begin to fuse.
- Fusion of two helium nuclei **does not produce a stable isotope**:  
 ${}^4\text{He} + {}^4\text{He} \rightarrow {}^8\text{Be}$  (lifetime  $\sim 10^{-16} - 10^{-17}$  s)

## Triple alpha process (at $10^8$ K)

Three helium nuclei can fuse simultaneously to produce carbon-12 (stable):

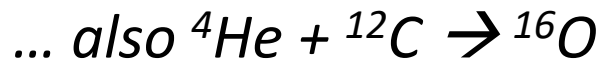


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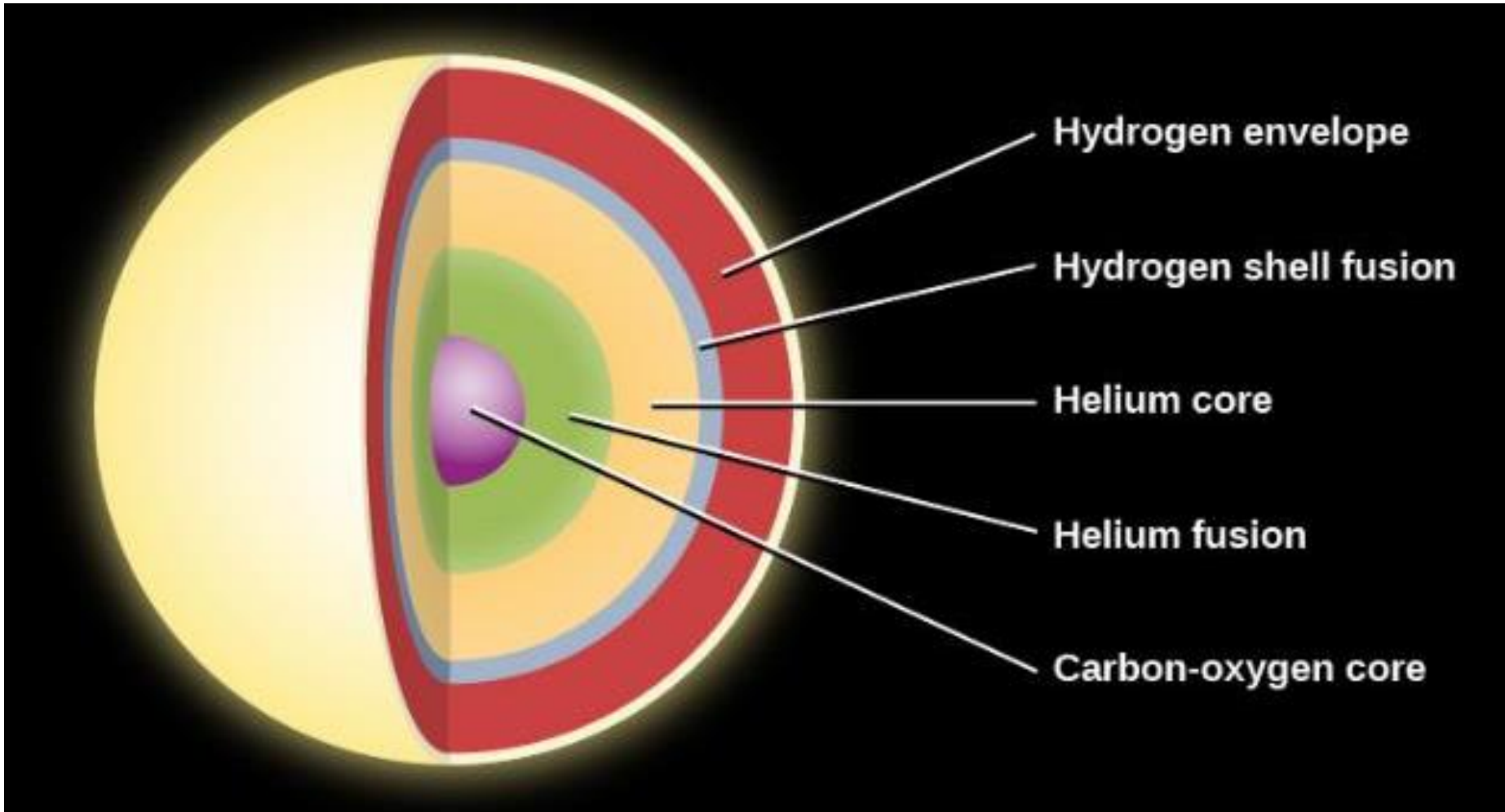
## Helium Flash (for Sun-like stars)

The fusion of helium into carbon happens very quickly (possibly in a few minutes).

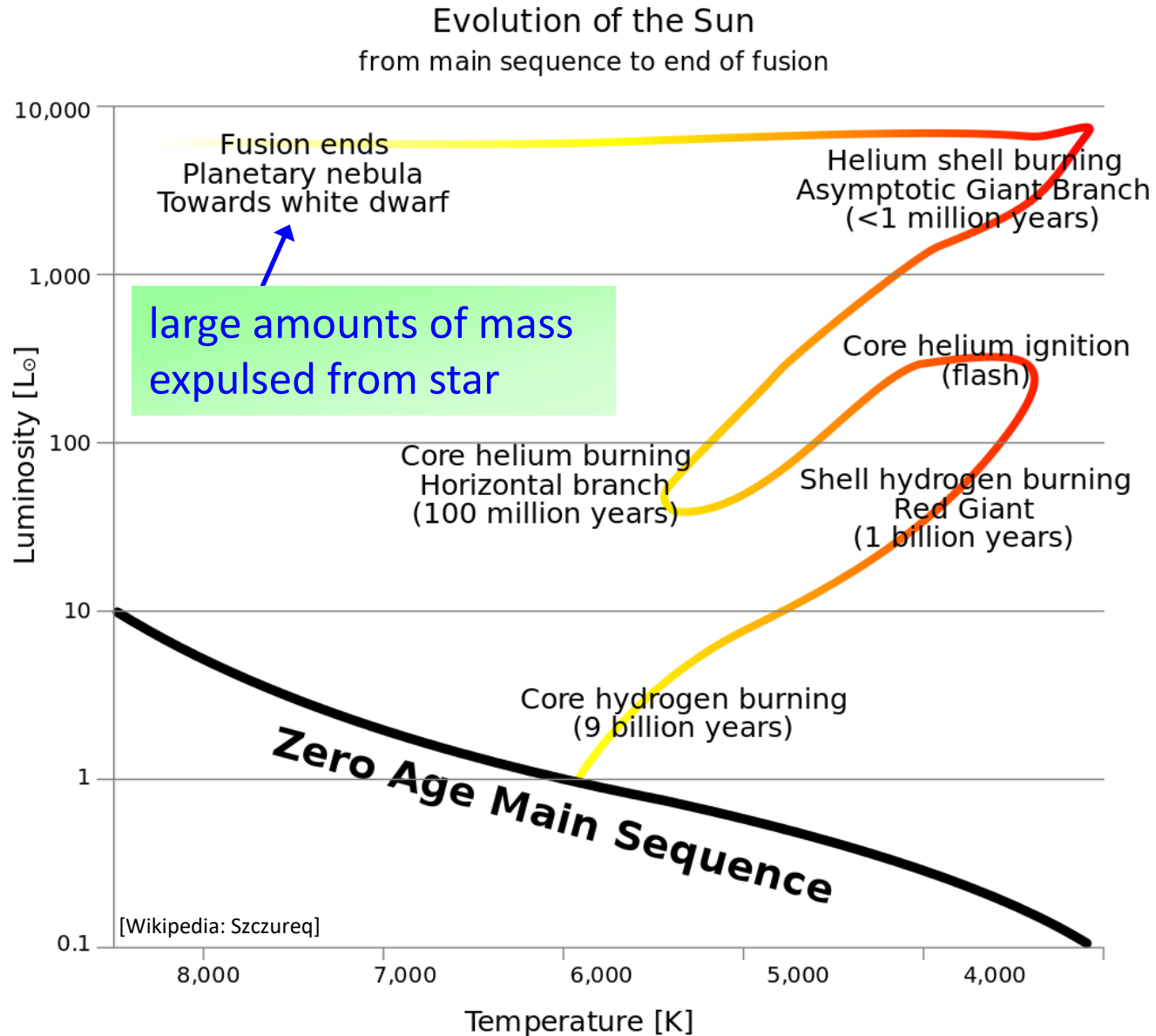




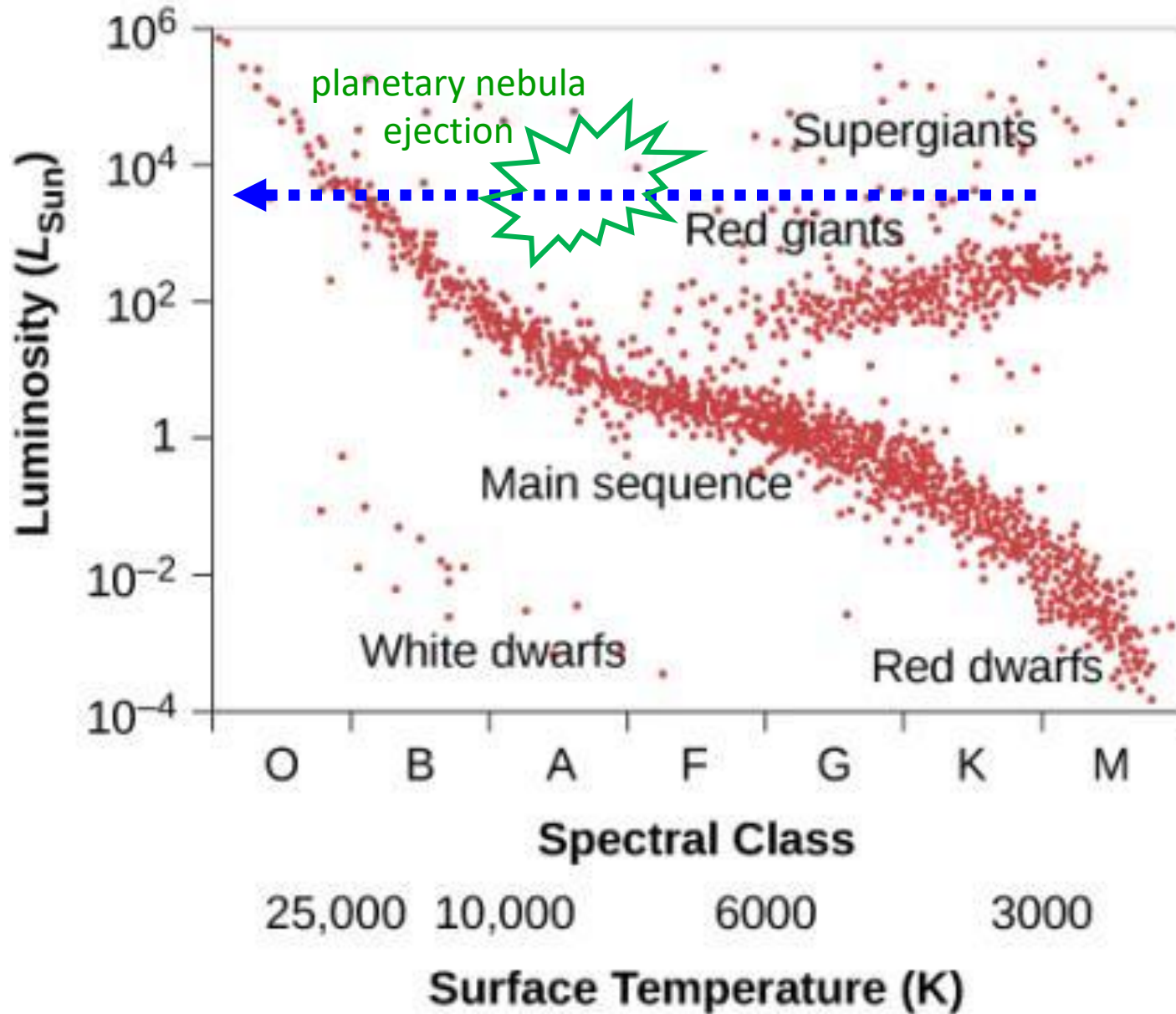
# Structure of Red Giant Star before “Death”



# Red Giant Evolution from Sun-like Star



# Planetary Nebula



# Mass Loss: Planetary Nebula

- Over the course of its red giant phase, a Sun-like star is expected to **shed roughly 50% of its mass**. Gas speed  $\sim$  20-30 km/s.
- This ejected mass becomes a **planetary nebula** with a **white dwarf** at its center.

*(note: planetary nebula has nothing to do with planets)*

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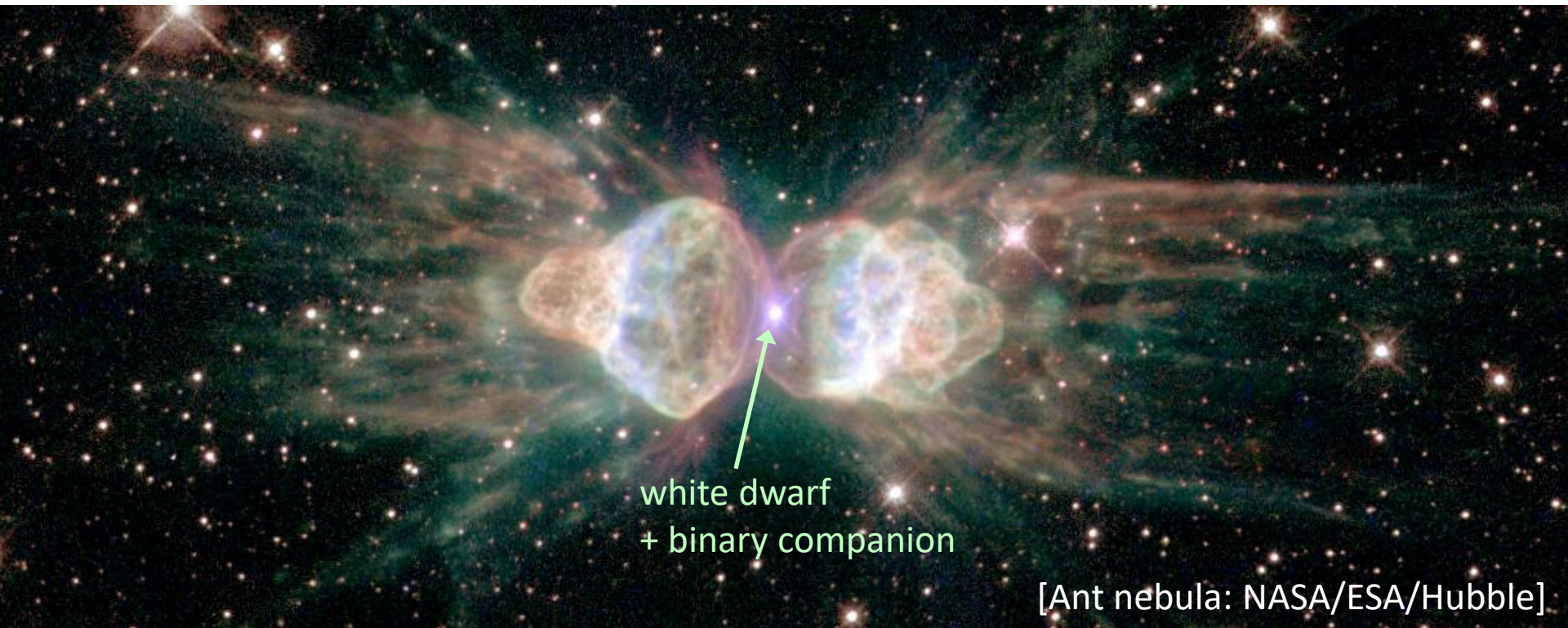
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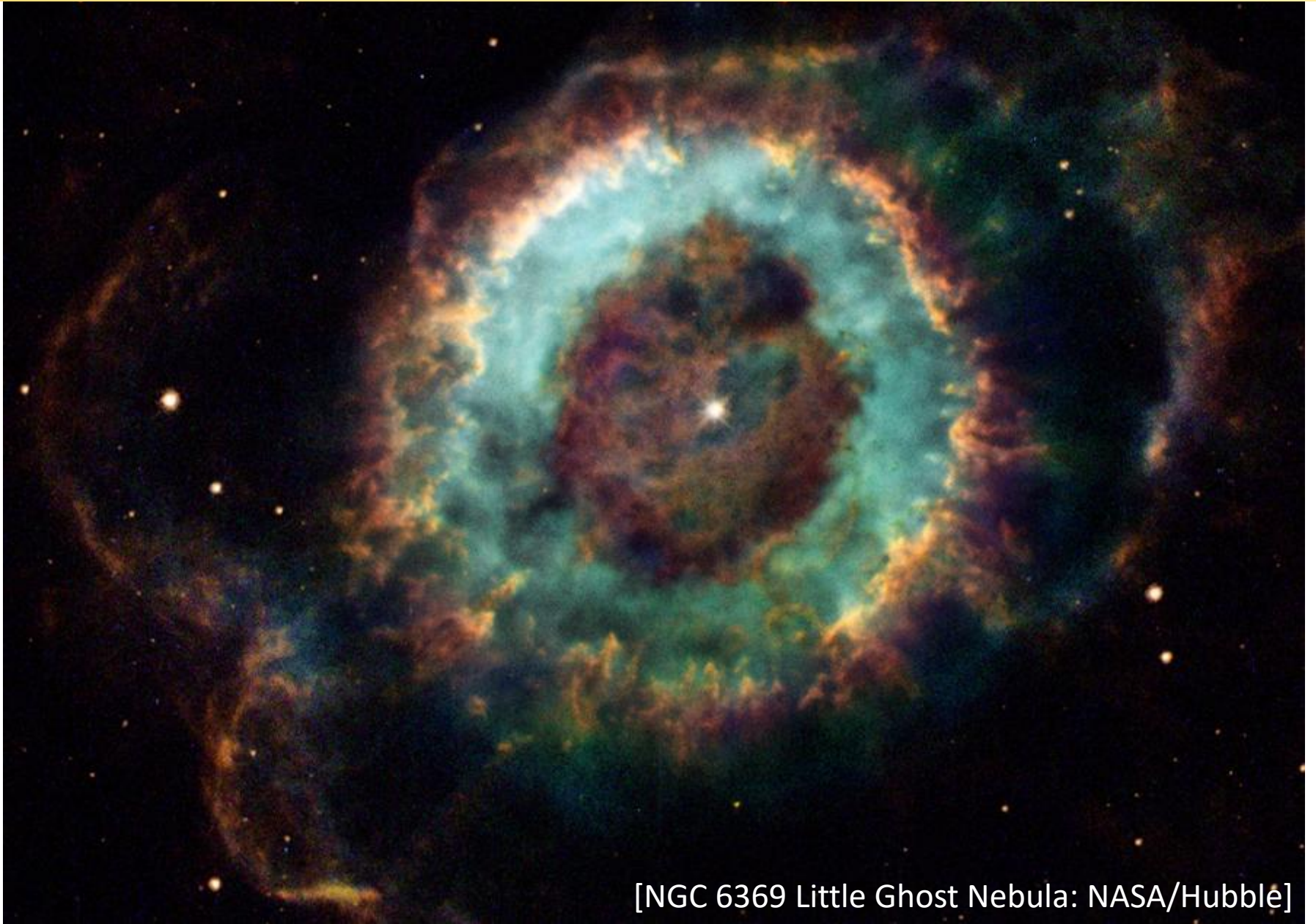
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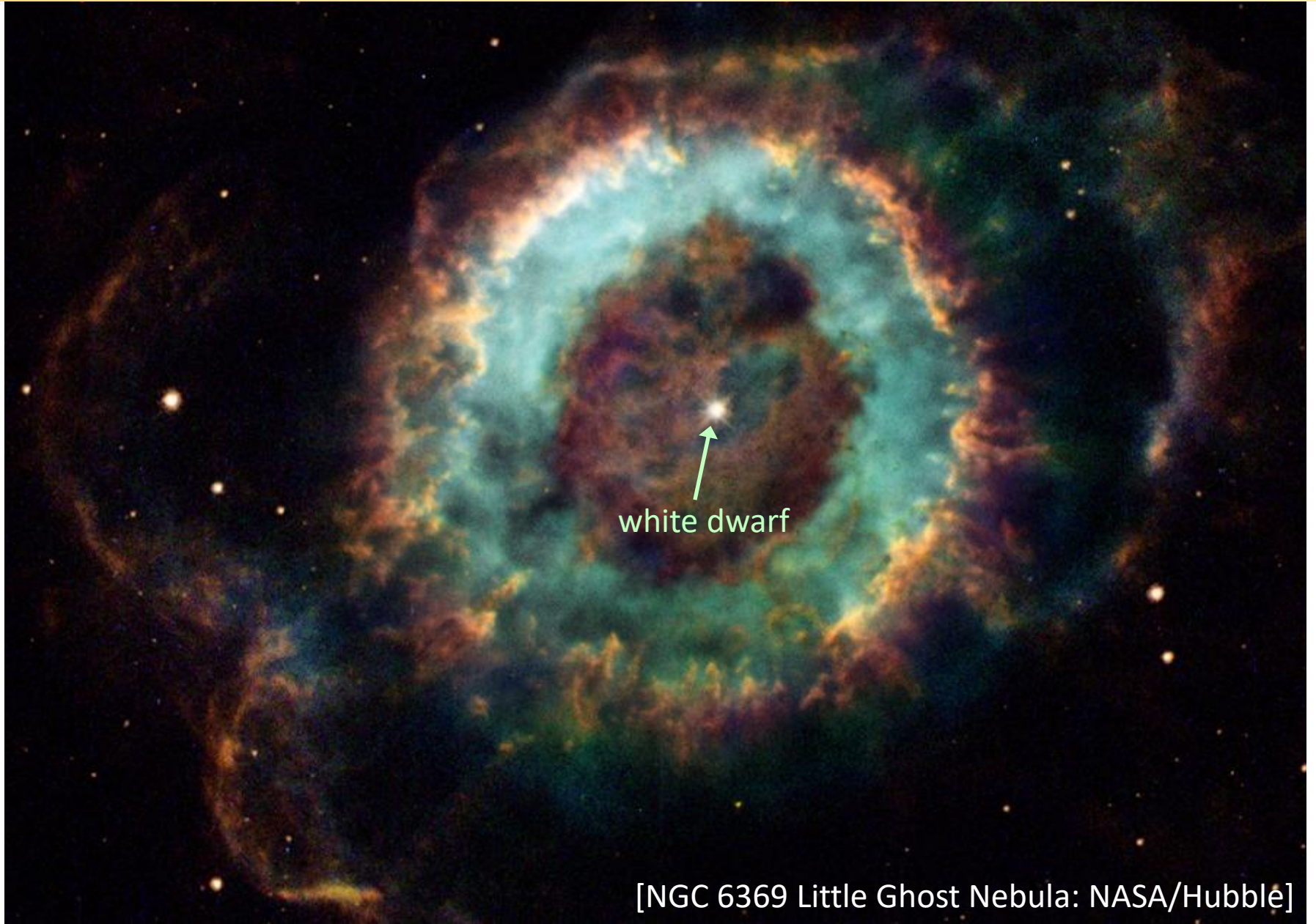


# Mass Loss: Planetary Nebula



[NGC 6369 Little Ghost Nebula: NASA/Hubble]

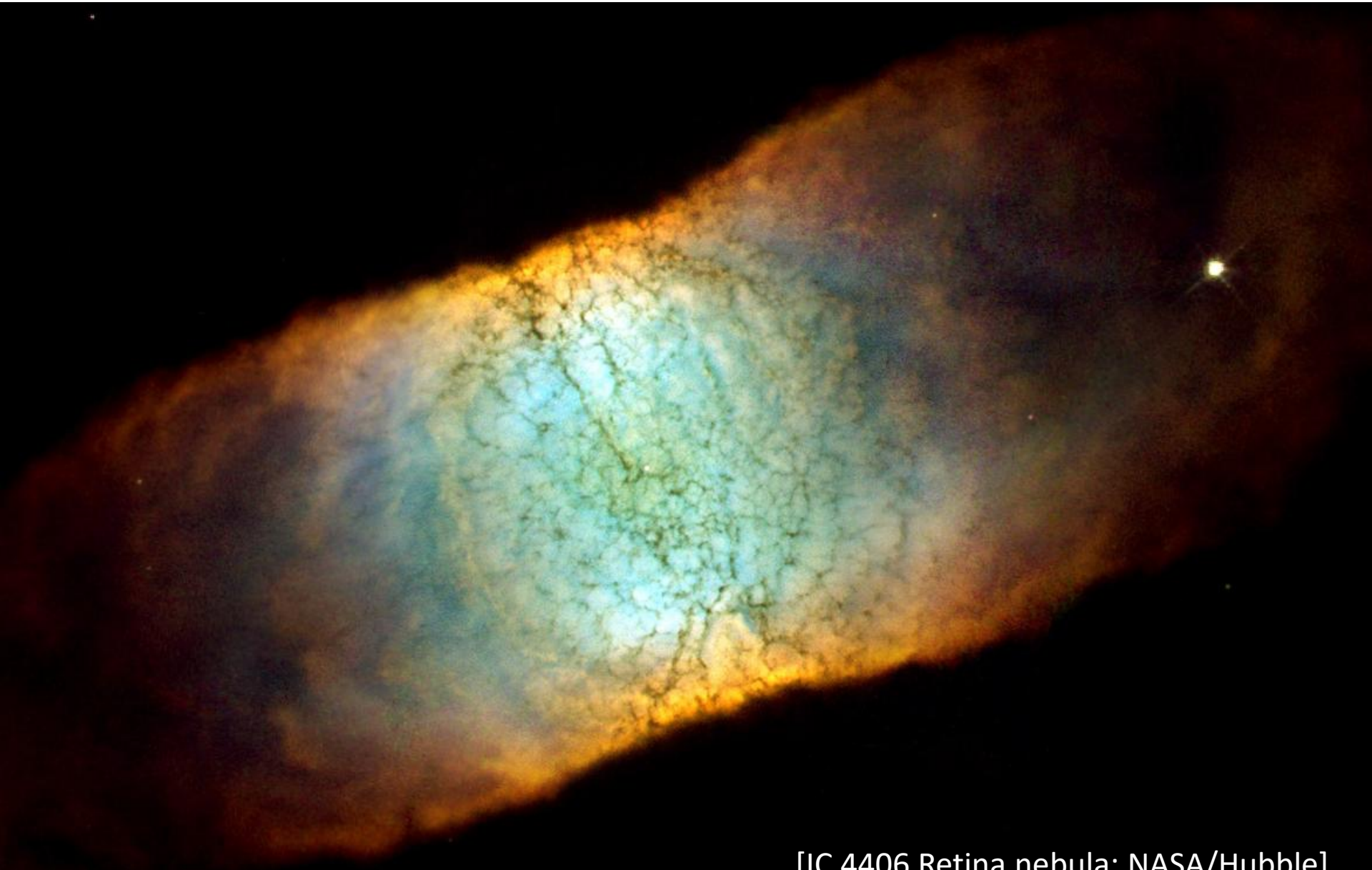
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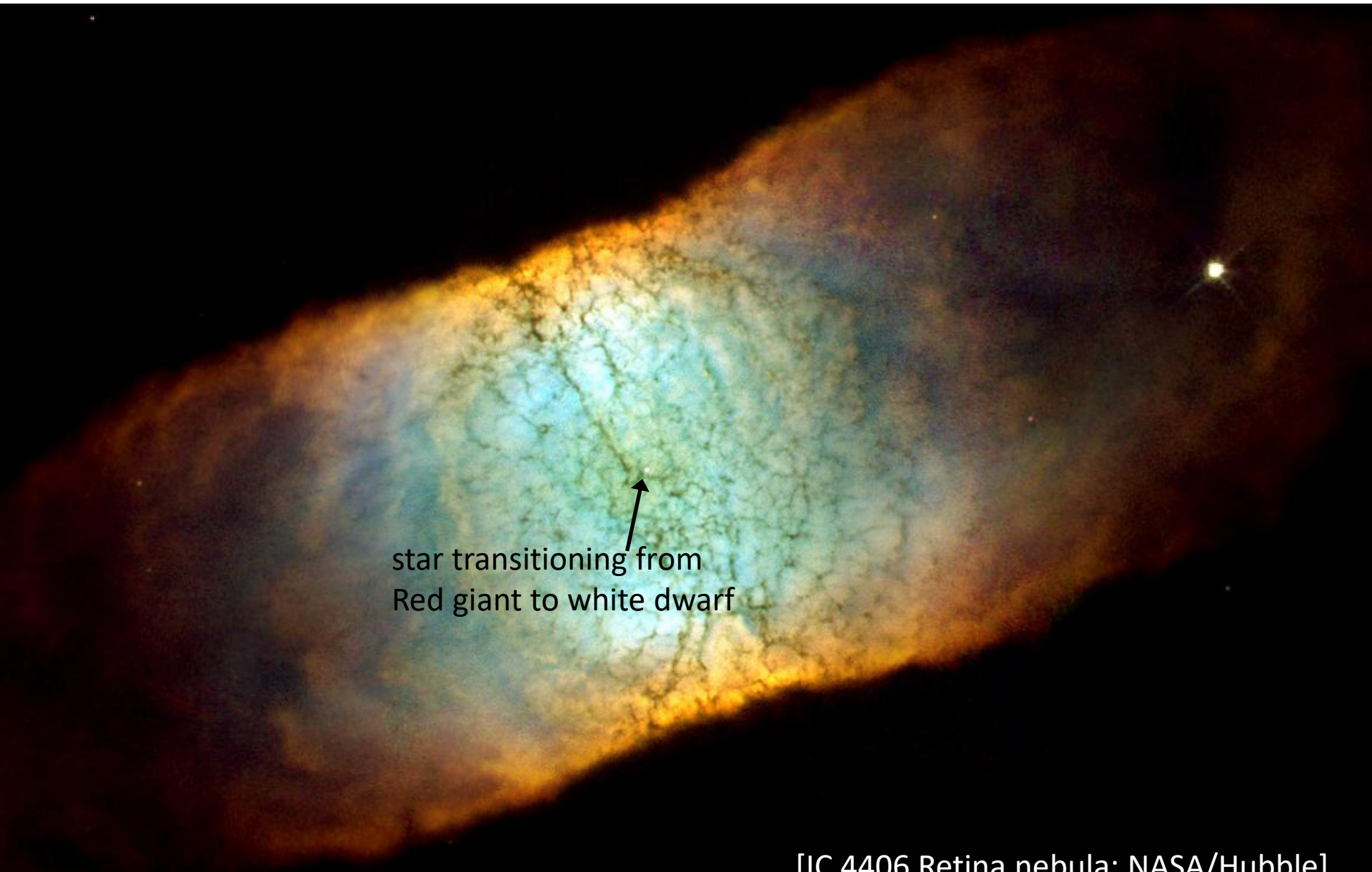


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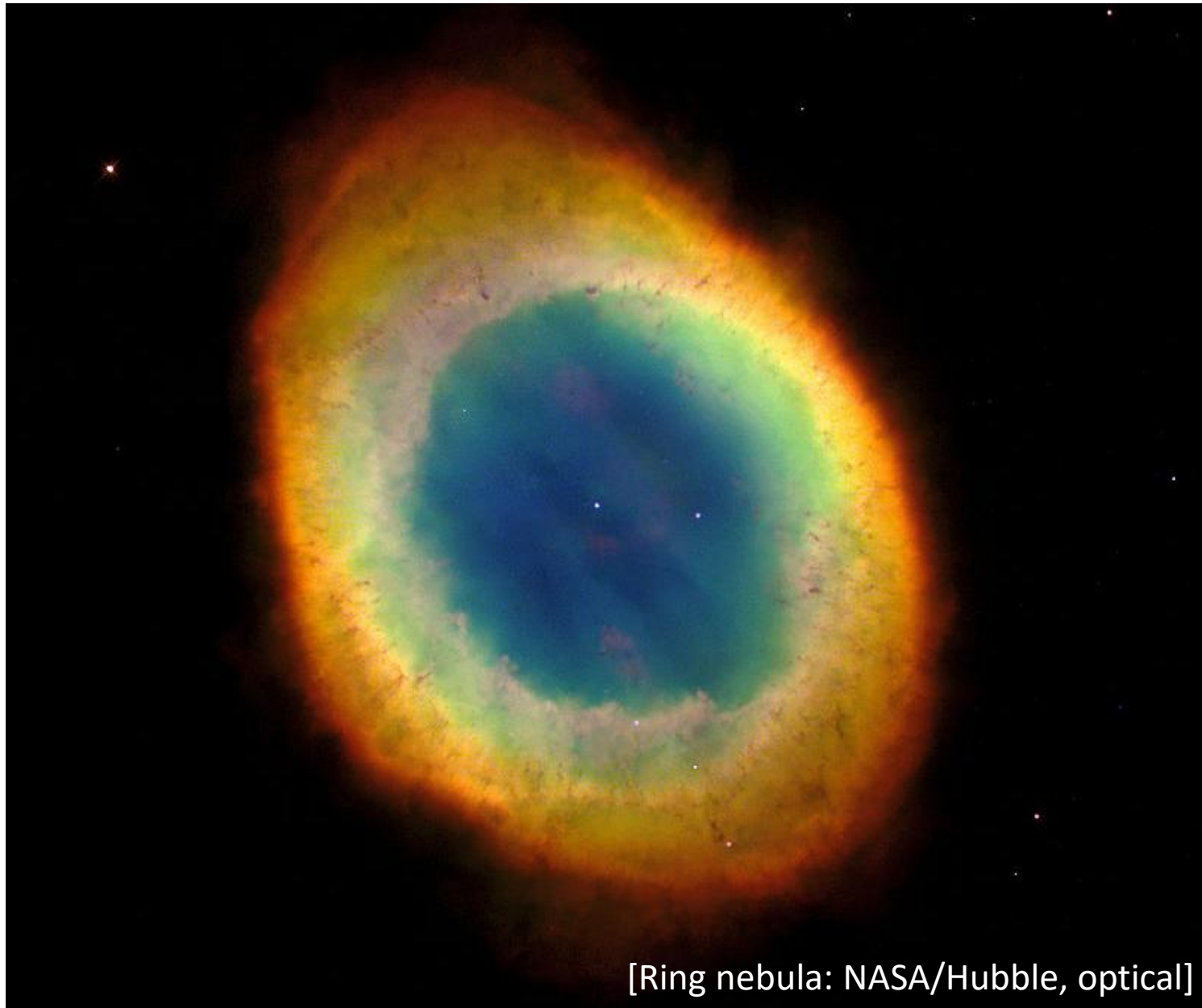
[IC 4406 Retina nebula: NASA/Hubble]

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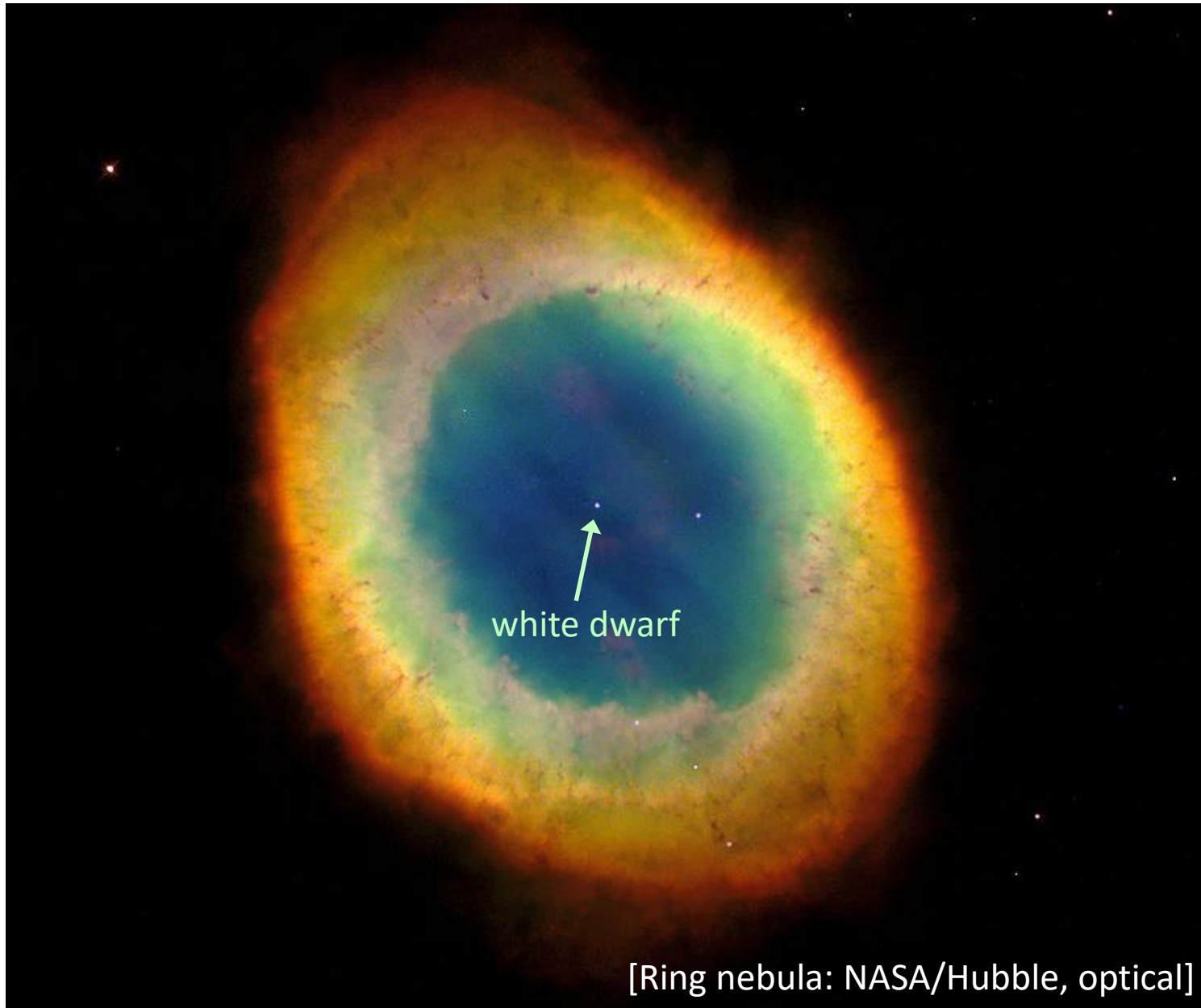
star transitioning from  
Red giant to white dwarf

# Mass Loss: Planetary Nebula



[Ring nebula: NASA/Hubble, optical]

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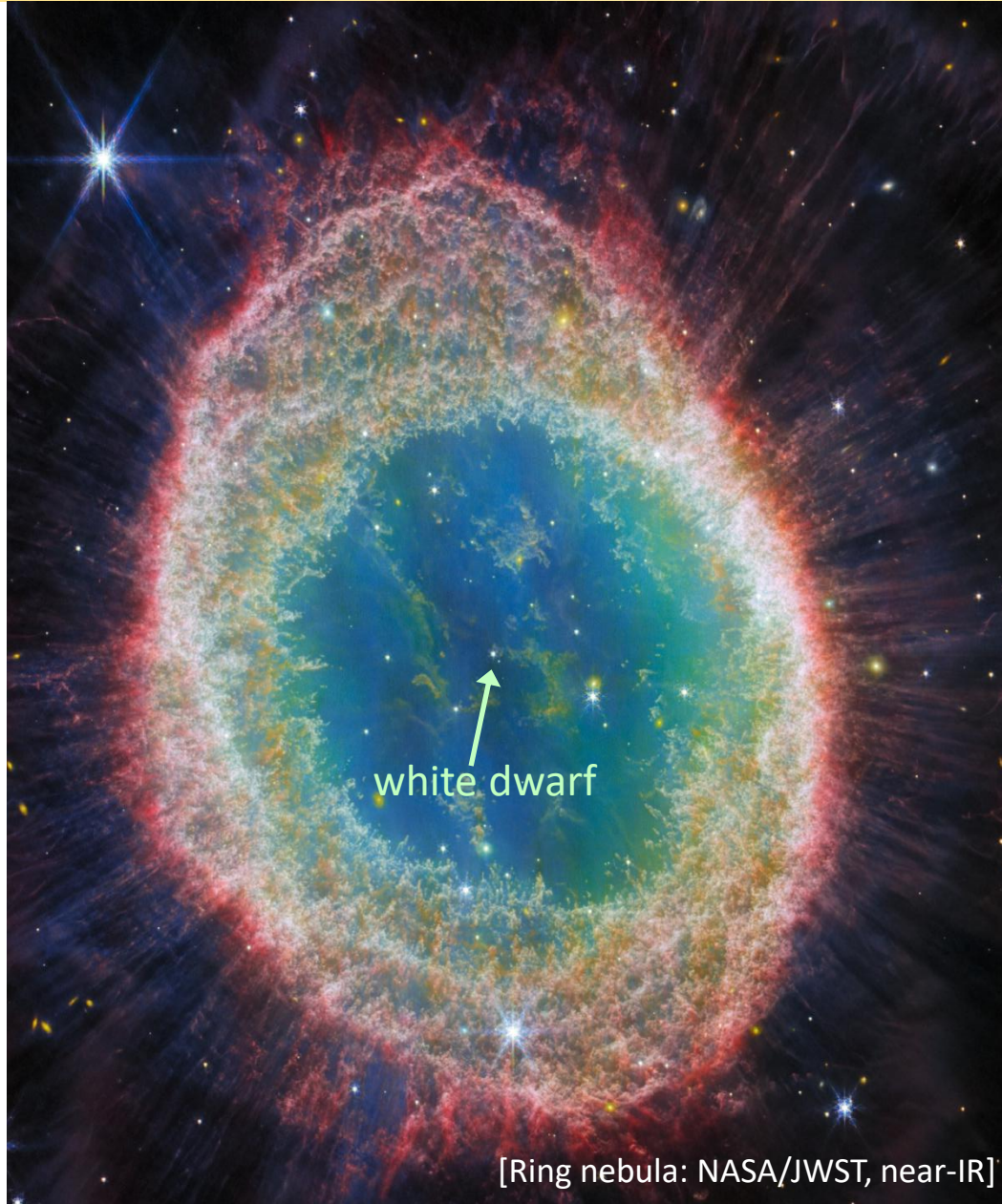
[Ring nebula: NASA/Hubble, optical]

# Mass Loss: Planetary Nebula



[Ring nebula: NASA/JWST, near-IR]

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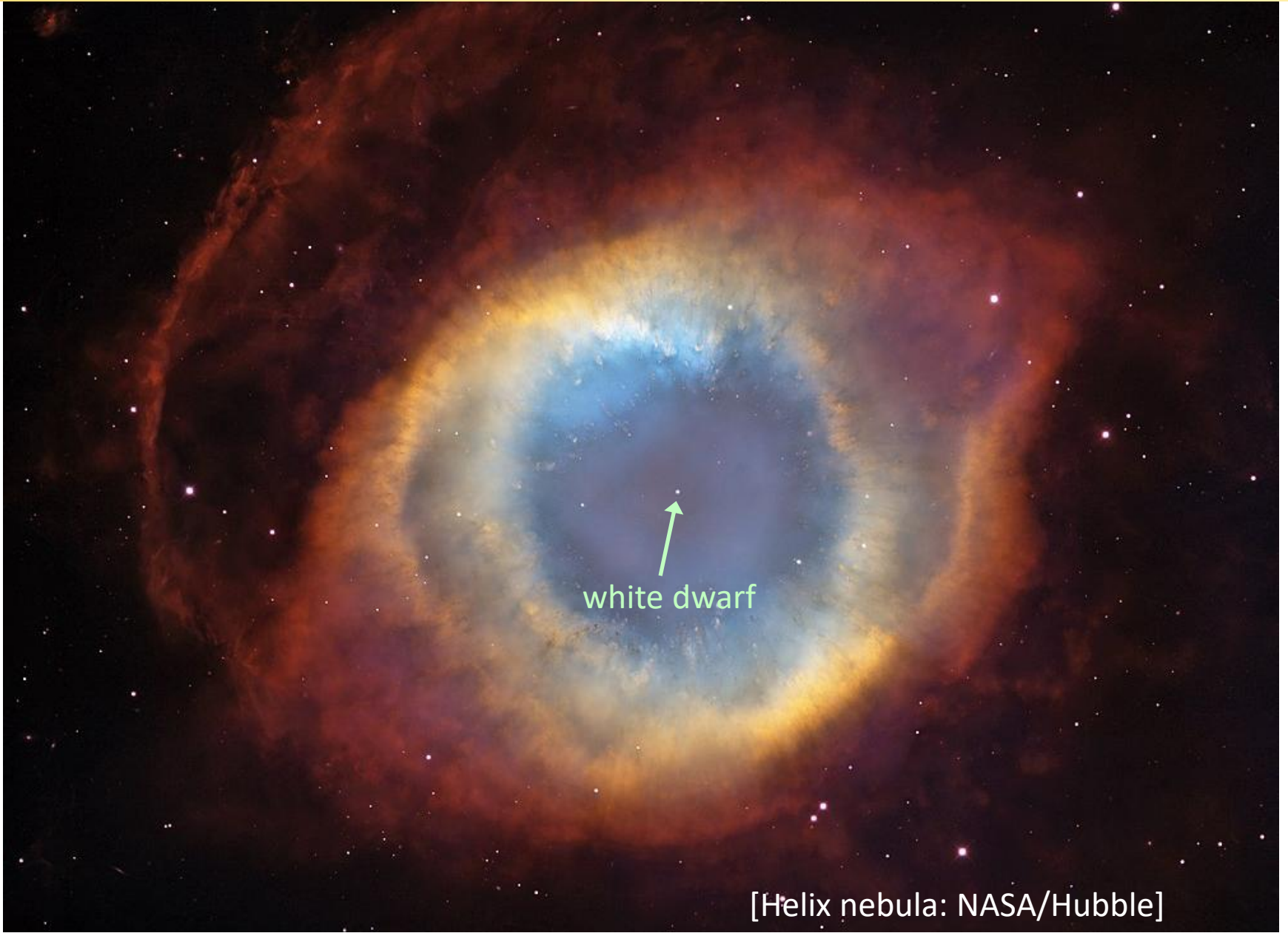
[Ring nebula: NASA/JWST, near-IR]

# Mass Loss: Planetary Nebula



[Helix nebula: NASA/Hubble]

# Mass Loss: Planetary Nebula



[Helix nebula: NASA/Hubble]



# Mass Loss: Planetary Nebula



[M2-9 Twin Jet / Butterfly Wings Nebula: ESA/Hubble]

# Mass Loss: Planetary Nebula



[NGC 6302 Butterfly Nebula: NASA/ESA/Hubble SM4 ERO Team]

# Mass Loss: Planetary Nebula



white dwarf or soon-to-be white dwarf  
(no binary companion)

$0.6 M_{\text{sun}}$ , 200,000 K [Szyszka et al, *Astrophys. J.* 707, L32 (2009)]

[NGC 6302 Butterfly Nebula: NASA/ESA/Hubble SM4 ERO Team]

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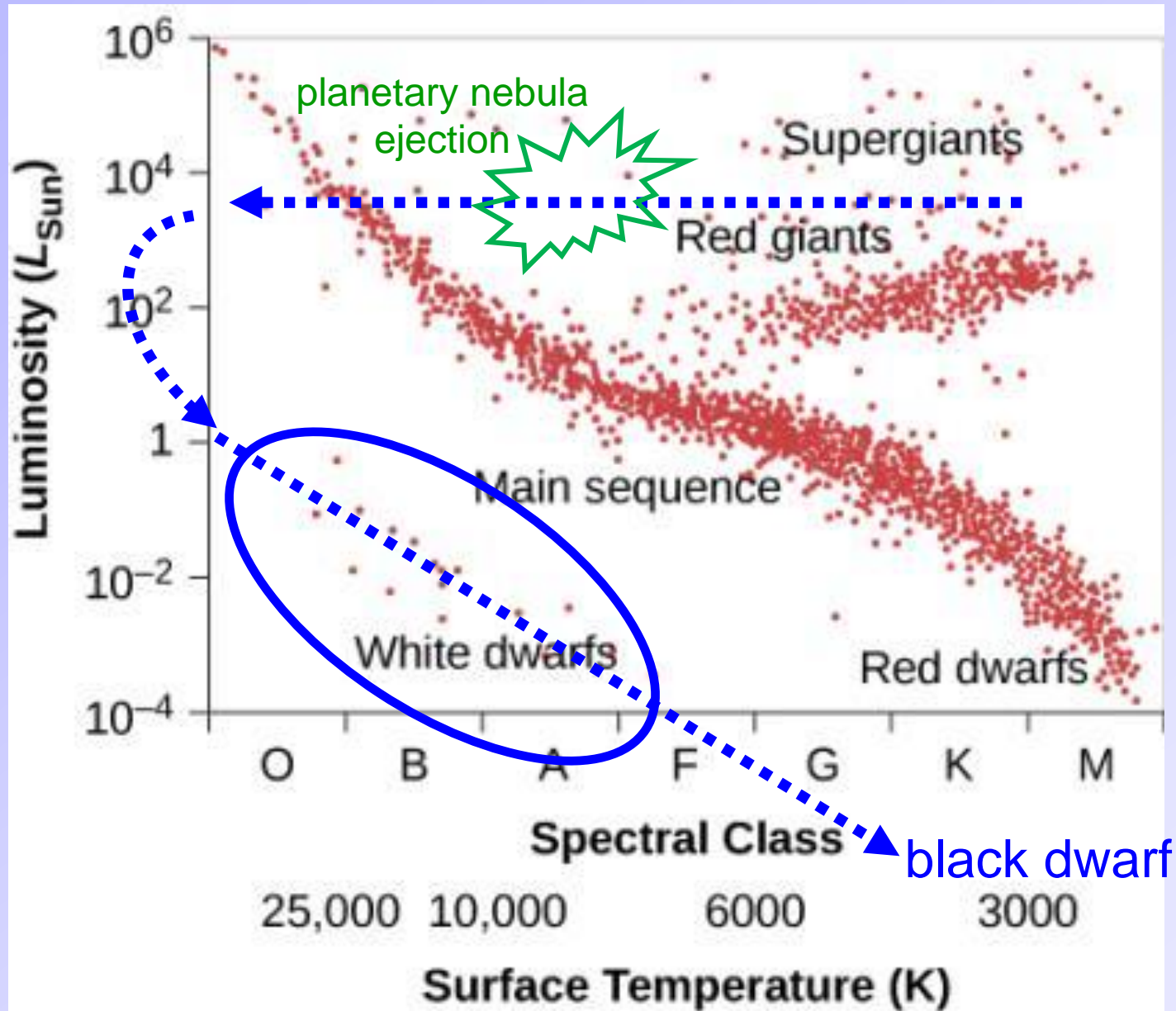


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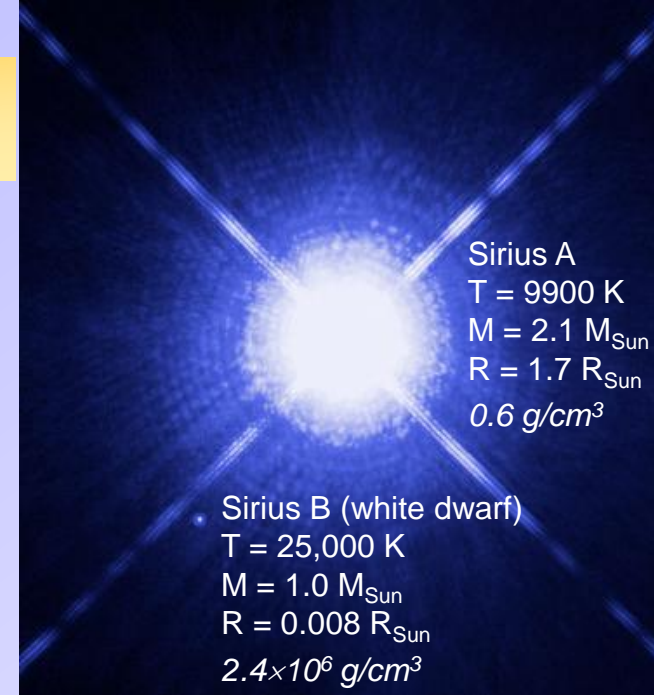


**PolleEv Quiz: [PolleEv.com/sethaubin](http://PolleEv.com/sethaubin)**

# White Dwarfs

## White dwarf

- “Ember” of dead star.
- Does not produce any energy of its own.
  - No fusion
- Starts out “white hot” and cools down to a black dwarf.
- Cools by emitting blackbody radiation.
- Heavier white dwarfs are smaller !!!

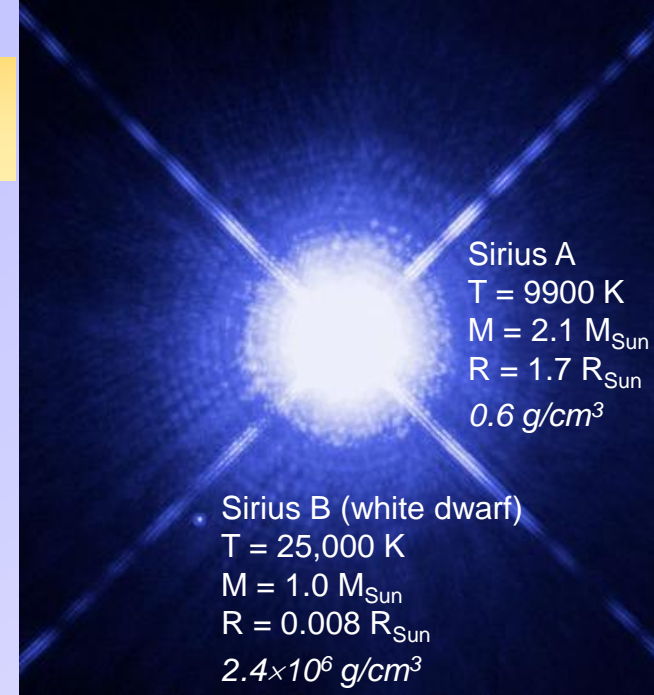
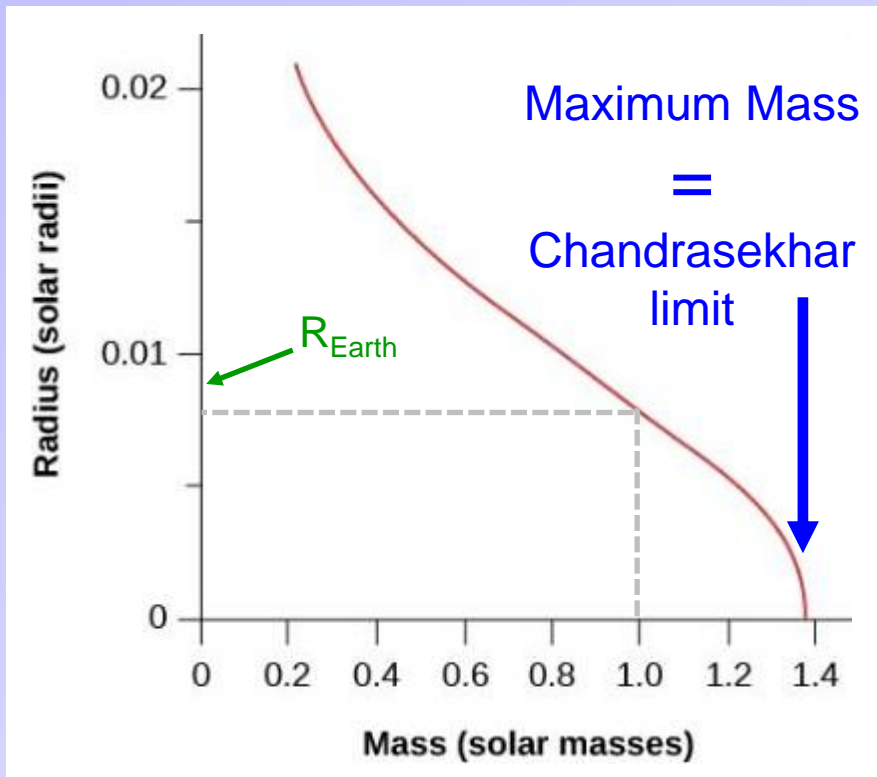


[NASA, ESA, H. Bond (STScI), and M. Barstow (University of Leicester)]

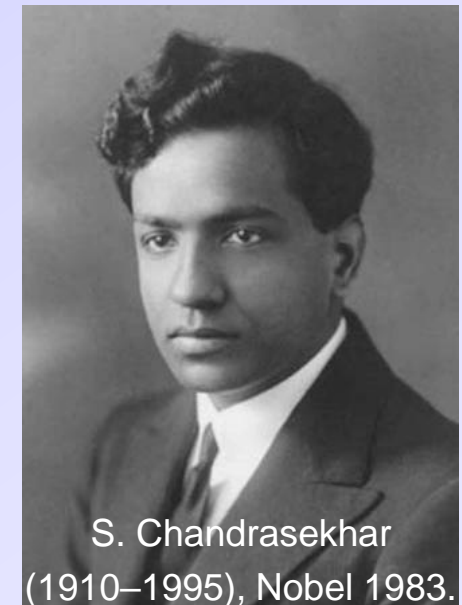
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# White Dwarfs

- A white dwarf is dense enough that ***gravity & pressure are strong enough to overwhelm the electric repulsion*** between nuclei and electrons, but ...
- Gravity is counteracted by quantum “**Pauli pressure.**”
- the Pauli exclusion principle for electrons: you cannot have more than one electron per quantum state (location or velocity).
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## White dwarf crystallization

**Below ~ 4,000 K**, the electric force between nuclei is strong enough to make an ordered arrangement of nuclei, i.e. “**nuclear crystal.**”

→ The **core** of the white dwarf **crystallizes.**

→ Some asteroseismology evidence.

